

Anisotropic assembly bias

in theory, simulations and data

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with Will Percival & Neal Dalal

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[arXiv:1906.11823](https://arxiv.org/abs/1906.11823)

[arXiv:2004.07240](https://arxiv.org/abs/2004.07240)

Outline

Rzdp@<zb^

, ^Sbzdp& P- b - ssC\ 4Y%4S s S^ sS \ ~Yzb^s

, ^Sbzdp& L- Y†%oo ssC\ 4Y%4S s S^ 3arr s- \ eY

; b^sC \ ~C^<Cs fHq? Br Rg. r~\ \ - q%oo

Overview

„ CP- fC- fCq%os~<<CsshY ; ? [\ b@CY@Cs<q4S'L zPC
} ^SfCpC

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; bs\ bYLS- YS'Hdq\ - zS^ <b^z- S^C@S^ @S CqC^z b4sCqf- 4Cs

Overview

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RPb\ bLC^SSCs @b\ S- ^z sb~qC bHS'Hbq\ - zSb^>\ - S^Y%o
zPp~LP | QbS^z sz zSSzSs bH• ~<z~ zSb^s

Overview

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3~z X qCCq<- YCrzq<z~qCS {? CteCzC@zb ~YS\ - zCY%P- fC
\ bqC<b^szq S'S'L eb.Cq

Overview

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\ bCq<b^szq S'S'L eb.Cq
} e<b\ S'L L- Y†%qC@PSH s~qfC%sof? Br R>B~<Y@g..SYqC <P
~^eq<C@C^zC@eq<SSb^

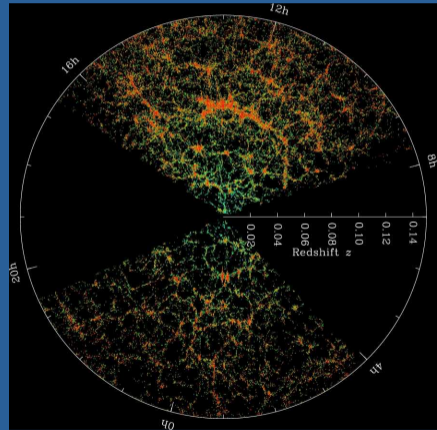
Large-scale structure

a. $f_{\text{CMB}} \approx 10^{-5}$

$$\langle \delta^2 \rangle = \langle \delta^2 \rangle / \langle \delta^2 \rangle = 1$$

b. $\langle \delta^2 \rangle \approx 10^{-5}$

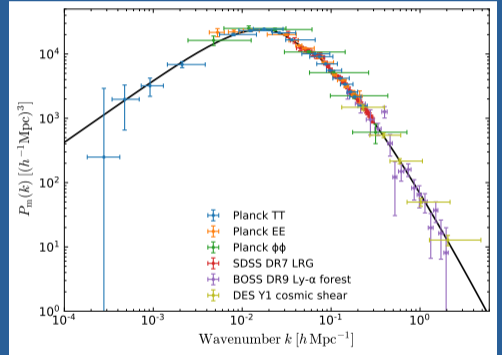
$$d \langle \delta^2 \rangle / h \langle \delta^2 \rangle = \langle \delta^2 \rangle$$



SDSS

Large-scale structure

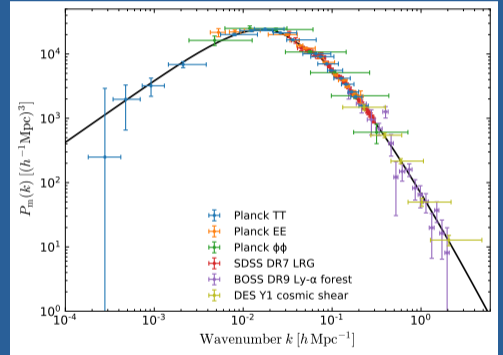
$$\begin{aligned}
 & \delta(\mathbf{k}) = \delta(\mathbf{k}) / \delta(\mathbf{k}) \\
 & d_{\text{vir}}(\mathbf{W}, \mathbf{W}) / h \delta(\mathbf{W}) \delta(\mathbf{W}) \\
 & ; b_{\text{vir}} \delta(\mathbf{W}) - Y d_{\text{vir}} \delta(\mathbf{W}) = d_{\text{vir}}(\mathbf{W})
 \end{aligned}$$



Planck, 2018

Large-scale structure

$\delta(k) = \delta(k) / \delta(k) = 1$
 $d_{\text{vir}}(W) / h d_{\text{vir}}(W) = d_{\text{vir}}(W)$
 $\text{Obs. } \delta(k) \sim k^{-1} P(k) \sim k^{-1} \delta(k)$
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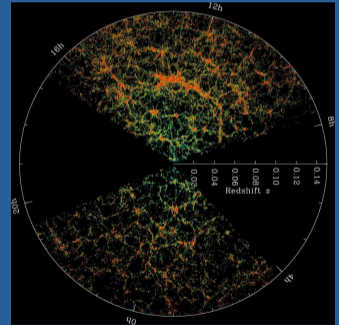


Planck, 2018

Linear bias and redshift-space distortions

Galaxies, halos, voids, 21cm, Ly forest ... all biased tracers of matter in real space, observed in redshift-space

$$d_L(W) = \frac{4\pi}{3} W^3 \quad d_L(W) = \frac{4\pi}{3} d^3(W)$$

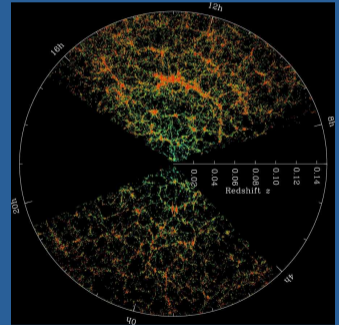


SDSS

Linear bias and redshift-space distortions

Galaxies, halos, voids, 21cm, Ly forest ... all biased tracers of matter in real space, observed in redshift-space

$$d_L(W) = \frac{d_L}{W} \left(\frac{dL}{dV} \right) \quad d_L(W) = \frac{d_L}{W} \left(\frac{dL}{dV} \right)$$



SDSS

Linear bias and redshift-space distortions

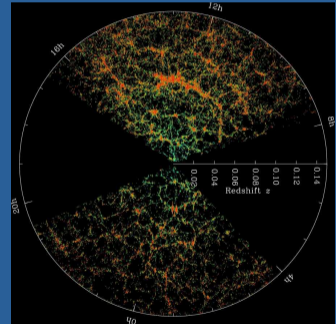
Galaxies, halos, voids, 21cm, Ly forest ... all biased tracers of matter in real space, observed in redshift-space

$$d_L(z) = \int_0^z \frac{c dt}{a(t)} \quad d_L(z) = \int_0^z \frac{c da}{a^2 H(a)}$$

$$d_L(z) = \frac{c}{H_0} \int_0^z \frac{dz'}{E(z')}$$

$$d_L(z) = \frac{c}{H_0} \int_0^z \frac{dz'}{E(z')}$$

- massive objects more biased
- objects more biased earlier



SDSS

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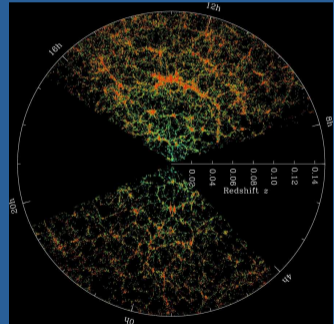
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- massive objects more biased
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$$d_L(W) = \frac{d_L}{W} \quad d_L(W) = \frac{d_L}{W}$$

$$d_L(W) = (1 + H^2) \quad d_L(W) = \frac{d_L}{W}$$



SDSS

Linear bias and redshift-space distortions

Galaxies, halos, voids, 21cm, Ly forest ... all biased tracers of matter in real space, observed in redshift-space

$$d_L(W) = \frac{d_L}{H} \left(\frac{dL}{dz} \right) \quad d_L(W) = \frac{d_L}{H} d \left(\frac{dL}{dz} \right)$$

$$\frac{dL}{dz} = \frac{c}{H(z)} \left(1 + \frac{v_r}{c} \right)$$

$$\frac{dL}{dz} = \frac{c}{H(z)} \left(1 + \frac{v_r}{c} \right) \left(1 + \frac{v_t}{c} \right)$$

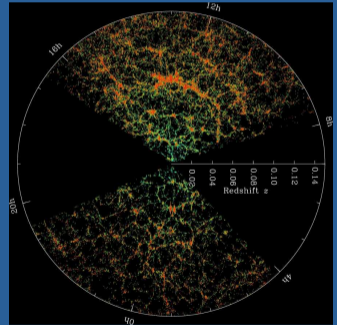
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SDSS

Galaxy power spectrum in redshift-space

Linear theory: $d_L^s(W) = (4_L + H^2)^2 d_L(W)$

Use Legendre expansion into multipoles:

$$d_L^s(W) = \sum_{\ell=0}^{\infty} \frac{2\ell+1}{2} d_L^s(W) L_{\ell}(\mu)$$

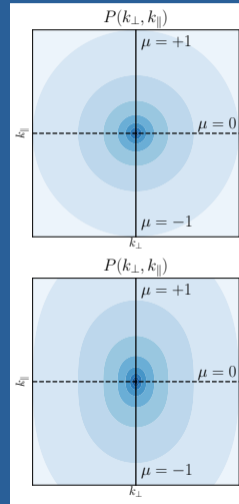
$$d_0(W) = 4_L + \frac{2}{3} H^2 4_L + \frac{1}{5} H^2 d_L(W)$$

$$d_2(W) = \frac{4}{3} 4_L H^2 + \frac{4}{7} H^2 d_L(W)$$

Measuring d_0 . d_2 gives 4_L . H

Note quadrupole $d_2 \neq 0$

In real-space $d_2 = 0$



Galaxy power spectrum in redshift-space

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Use Legendre expansion into multipoles:

$$d_L^s(W) = \sum_{l=0}^{\infty} \frac{2^l + 1}{2} P_l(\cos\theta) d_L^s(W) L_l(\cos\theta)$$

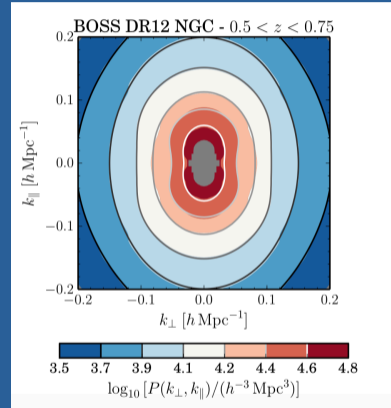
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Measuring d_0 . d_2 gives 4_L . H

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Alam+2016

Growth rate H

One of the key parameters

$$H = \frac{d \ln a}{dt}$$

$$K_p = \frac{f\sigma_8}{h} = H \frac{m(\zeta)^{0.55}}{h}$$

$$R_{\text{eff}} = \frac{f\sigma_8}{h}$$

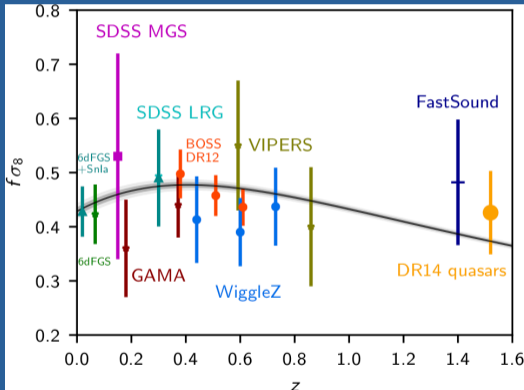
- Testing Gravity
- Constraining neutrino masses
- Testing dark energy models
- ...

$$\sigma_8 \approx 5 \text{ } h$$

$$G = \frac{f\sigma_8}{h} \approx 1 \text{ } h$$

$$f\sigma_8 \approx 1 \text{ } h$$

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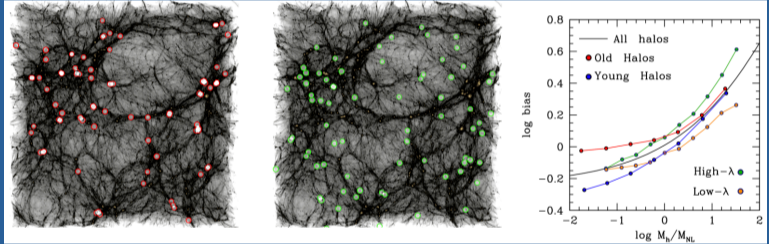


Planck, 2018

Assembly bias

Bias depends on other $s \ll Yq$ properties, for fixed halo mass and redshift

- Formation history
- Age
- Spin
- Concentration
- Shape ...

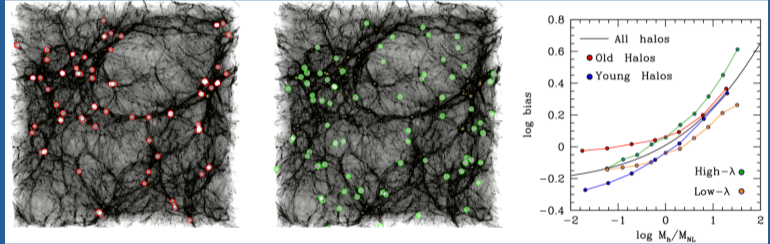


Wechsler+, 2018

Assembly bias

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- Formation history
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Wechsler+, 2018

? $G \ll z \ll S \ll s \ll Y \ll b \ll s \ll b \ll f \ll S \ll L \ll G \ll S \ll C \ll S \ll @ \ll z$

Non-scalar bias

$$\mathbb{E}[\hat{\beta}] = \beta + (X'X)^{-1}X' \mathbb{E}[e] \quad \text{where } \mathbb{E}[e] = (X'X)^{-1}X' \mathbb{E}[e] + H^2 \quad \text{and } H = (X'X)^{-1}X'X$$

Non-scalar bias

$$\begin{aligned}
 & \text{, } z z P C \mathcal{B}^{\wedge} G q Y f C Y S^{\wedge} \setminus \dots C \sim s \sim \mathbb{W} / \infty s s \sim \setminus G = L(\mathbf{W} = (4_L + H^2) \setminus (W \\
 & y P G C s - ^{\wedge} b z P G q z C q \setminus \mathcal{B}^{\wedge} G q S^{\wedge} \setminus \quad z q < C Y s s e - q z b H P C z S @ Y' C Y @ \\
 & s_S(\ddagger) = (r \ s' \ U' \ ^2 \quad s/3) \setminus (\ddagger) \quad () \quad s_S(\mathbf{W} = (W W W \quad s/3) \setminus (\mathbf{W}
 \end{aligned}$$

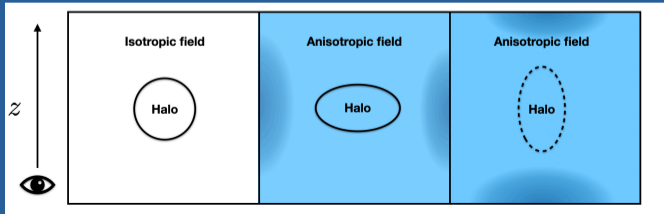
Non-scalar bias

$$\begin{aligned}
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 & y P G C s - ^b z P G q z C q \setminus \mathcal{B} \wedge G q S \setminus \quad z q < C Y s s e - q z b H P C z S @ Y' C @ = \\
 & s_S(\ddagger) = (r \ s' \ U' \ ^2 \quad s' / 3) \setminus (\ddagger) \quad () \quad s_S(\mathbf{W} = (W W W \quad s' / 3) \setminus (\mathbf{W} \\
 & \quad L(\mathbf{W} = (4_L + H^2) \setminus (W + 4_S s_S(\mathbf{W}
 \end{aligned}$$

Non-scalar bias

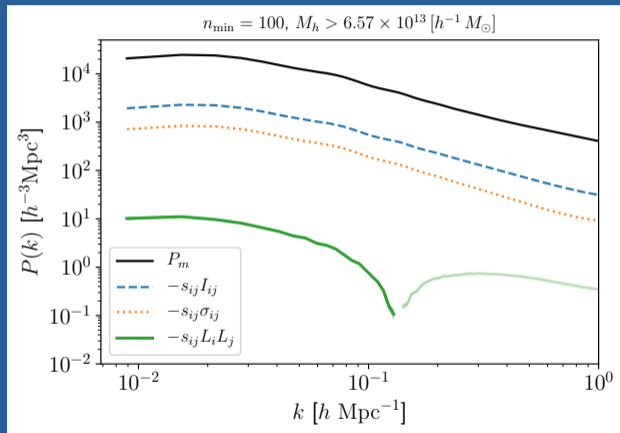
$$\begin{aligned}
 & \text{...} \\
 & \text{...} \\
 & s_{\text{eff}}(\mathbf{z}) = (r_{\text{eff}}^{-2} \text{...}) \\
 & L(\mathbf{W}) = (4L + H^2) \text{...}
 \end{aligned}$$

...
 – projected sizes, velocity dispersion & angular momentum



How correlated are halos & tidal field?

We use 1000 Quijote N-body sims (Villaescusa-Navarro+, 2019) to measure cross-correlations



, $a_j | \mathcal{E}_-$

Non-scalar bias

$$L(\mathbf{W}) = (4_L + H^2) \setminus (W + 4_{SS} s_{SS}(\mathbf{W}))$$

$$s_{SS}(\mathbf{W}) = (r \ s \ u \ ^2 \ s/3) \setminus (\mathbf{W})$$

$a \ ^{\wedge} \Psi \ ^{\wedge} b \ ^{\wedge} Q \leftarrow Y \ q \ e \ p \ e \ C \ o \ f \ S \ s \leftarrow \ ^{\wedge} \leftarrow b \ o \ p \ Y \ z \ C \ .. \ S \ P \ z \ @ \ Y \ ' \ C \ @$
 – projected sizes, velocity dispersion & angular momentum

$$, \ \langle S \rangle \sim \mathcal{P} \cdot Y \ s \ ^{\circ} \ \backslash \ C \ z \ ^{\circ} \ \backslash \ 4_{\langle} \quad 4_{\langle}$$

$$L(W) = (4_L + H^2) \setminus (W + 4_{\langle} (\ ^2 \ 1/3) \setminus (W))$$

$$= 4_L \ 4/3 + (H + 4_{\langle}) \ ^2 \ \setminus (W)$$

Non-scalar bias

$$L(\mathbf{W}) = (4_L + H^2) \setminus (W + 4_S S_S(\mathbf{W}))$$

$$= 4_L \left(\frac{2}{3} + (H + 4_L)^2 \right) \setminus (W)$$

$a \wedge \mathbf{Y} \hat{=} b \wedge \mathbf{Q} \leftarrow \mathbf{Y} \mathbf{q} \mathbf{e} \mathbf{p} \mathbf{e} \mathbf{C} \mathbf{r} \mathbf{S} \leftarrow \wedge \leftarrow \mathbf{b} \mathbf{q} \mathbf{Y} \mathbf{z} \mathbf{C} \dots \mathbf{S} \mathbf{P} \mathbf{z} \mathbf{S} \mathbf{Y} \mathbf{C} \mathbf{e}$
 – projected sizes, velocity dispersion & angular momentum

$$L(W) = (4_L + H^2) \setminus (W + 4_S (\frac{2}{3} + 1/3)) \setminus (W)$$

$$= 4_L \left(\frac{2}{3} + (H + 4_L)^2 \right) \setminus (W)$$

$$S_S(\mathbf{z}) = (r_S \mathbf{z} \mathbf{U}^2 \quad S_S/3) \setminus (\mathbf{z}) \quad S_S(\mathbf{W}) = (W \mathbf{W} \mathbf{W} \quad S_S/3) \setminus (\mathbf{W})$$

Anisotropic assembly bias (AB)

$$\frac{s}{L} = (4_L + H^2) \setminus E) \quad \frac{s}{L} = (4_L \quad 4/3 + (H \setminus 4) \quad ^2) \setminus$$

Anisotropic assembly bias (AB)

$$\frac{s}{L} = (4L + H^2) \setminus E) \quad \frac{s}{L} = (4L - 4/3 + (H - 4)^2) \setminus$$

d-q \setminus Cq4 S zPC anisotropic assembly bias

Anisotropic assembly bias (AB)

$$\frac{s}{L} = (4_L + H^2) \setminus E) \quad \frac{s}{L} = (4_L \quad 4/3 + (H \setminus 4) \quad ^2) \setminus$$

d-q \setminus C \setminus C q 4 \setminus S \setminus zPC anisotropic assembly bias

rb~qCbH ^Sbzqpe%8^ zPC qG Yse <Ceb. CqseG<zq \setminus

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d-q \ CZCq 4 \$ zPC anisotropic assembly bias

rb~qCbH ^Sbzpe%8^ zPC qG Yse <Ceb.CqseGzq\

, @@SS^ - Ysb~qCbH ^Sbzpe%8^ zPC q@SPzIqe <C

Anisotropic assembly bias (AB)

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d-q \ CZCq 4 \$ zPC anisotropic assembly bias

rb~qCbH ^Sbzdpe%\$ zPC qG Yse <Ceb.CqseGzq\

, @@SS^ - Ysb~qCbH ^Sbzdpe%\$ zPC q@SPzqe <C

] bzC 4 \$ eGqCzY%@LC^Gq zC..SP HF

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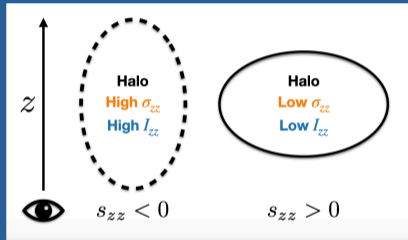
anisotropic assembly bias

$$4_i = 0 \quad \mathbb{H}$$

- Selection independent of halo orientation, e.g. projected size, velocity dispersion, angular momentum
- *but* if observed tracer and host halo randomly misaligned

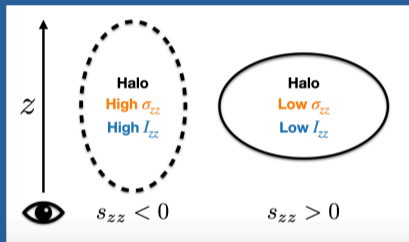
Halo selection based on tensor properties

Selection on radial halo extent & velocity dispersion σ_{zz} in real space



Halo selection based on tensor properties

Selection on radial halo extent & velocity dispersion σ_{zz} in real space

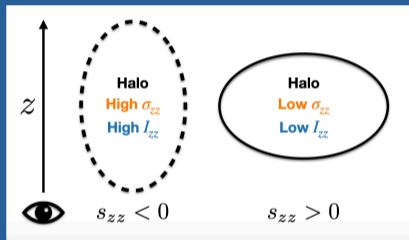


Real-space $d_2 = H \neq 0$

$d_2 \notin 0 ! \quad 4 \notin 0$

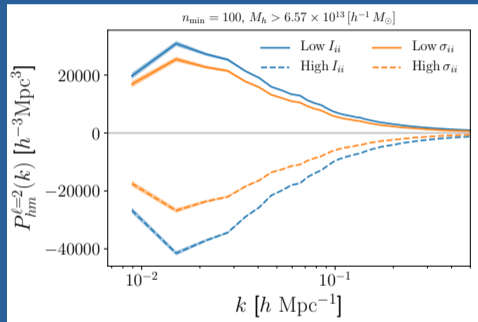
Halo selection based on tensor properties

Selection on radial halo extent & velocity dispersion $c?$ in real space



Real-space $d_2 = H \neq 0$

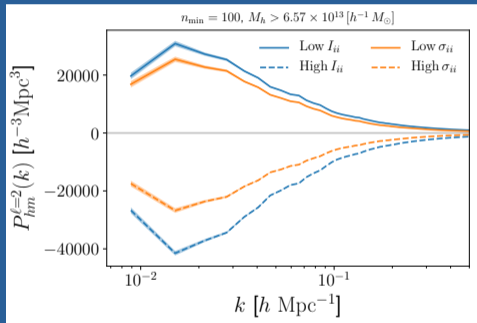
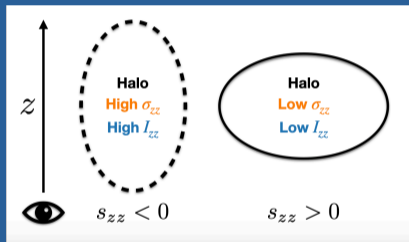
$d_2 \notin 0 ! \quad q \notin 0$



AO+2019

Halo selection based on tensor properties

Selection on radial halo extent & velocity dispersion σ_{zz} in real space



AO+2019

Real-space $d_2 = H \neq 0$

$d_2 \notin 0 ! \quad q_4 \notin 0$

Halos: $q_4 \quad 1 \quad 2$

Redshift-space $H \quad 0:7$

What about real galaxies?

When split on orientation dependent quantities, do galaxies show different clustering strength?

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When split on orientation dependent quantities, do galaxies show different clustering strength?

Baryon Oscillation Spectroscopic Survey

BOSS DR12 galaxy sample

10^6 L- Y†%q@sPSz

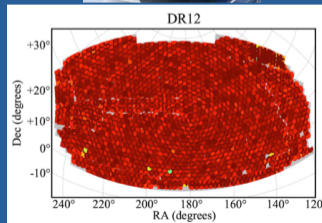
0:15 < < < 0:7

X- \ S b- s q@L- Y†SS> 4L 2

BYeZS- Y> [P 10¹³ [/P

K- Y†%s- \ eYS

- LOWZ (0:15 < < < 0:43)
- CMASS (0:43 < < < 0:7)



How do we look for AB?

Main idea – split on orientation (φ) / look for differences in anisotropy (χ)

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Subsamples matching $\hat{\chi}(\varphi)$ have matching H

How do we look for AB?

Main idea – split on orientation (θ) / look for differences in anisotropy (χ)

Subsamples matching $\hat{\chi}(\theta)$ have matching H

Subsamples can have different χ_L & χ_H

How do we look for AB?

Main idea – split on orientation (θ) ! look for differences in anisotropy (d_2)

Subsamples matching θ have matching H

Subsamples can have different d_1 & d_2

Find subsamples matching d_1 & θ !

Mismatch d_2 ! evidence $d_2 \neq 0$

How do we look for AB?

Main idea – split on orientation (θ) ! look for differences in anisotropy (d_2)

Subsamples matching θ have matching H

Subsamples can have different d_1 & d_2

Find subsamples matching d_1 & θ !

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Need orientation dependent gal. property

How do we look for AB?

Main idea – split on orientation (θ) ! look for differences in anisotropy (d_2)

Subsamples matching θ have matching H

Subsamples can have different d_1 & d_2

Find subsamples matching d_1 & θ !

Mismatch d_2 ! evidence $d_2 \neq 0$

Need orientation dependent gal. property

Galaxy Properties from Portsmouth Group

- velocity dispersion σ (1D)
- stellar mass $[M_\odot]$

How do we look for AB?

Main idea – split on orientation (β) ! look for differences in anisotropy (β)

Subsamples matching β have matching H

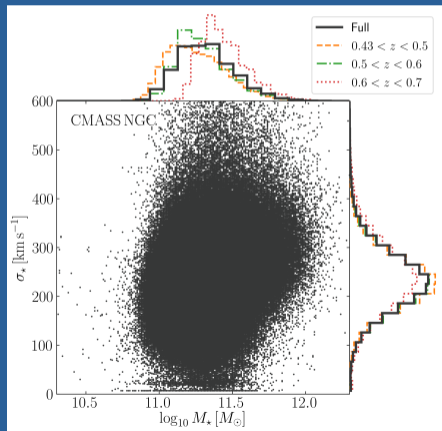
Subsamples can have different β & β

Find subsamples matching d_0 & β !

Mismatch d_2 ! evidence $\beta \notin 0$

Galaxy Properties from Portsmouth Group

- velocity dispersion β (1D)
- stellar mass β



How do we look for AB?

Main idea – split on orientation (β) ! look for differences in anisotropy (β)

Subsamples matching β have matching H

Subsamples can have different β & β

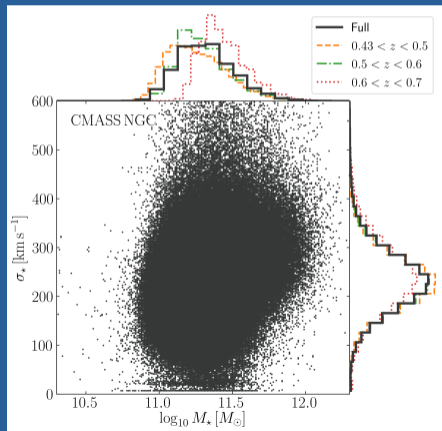
Find subsamples matching d_0 & β !

Mismatch d_2 ! evidence $\beta \notin 0$

Galaxy Properties from Portsmouth Group

- velocity dispersion β (1D)
- stellar mass $[\beta$

Split on $\beta =$ split on orientation & galaxy mass ($\beta / [\beta$)



How do we look for AB?

Main idea – split on orientation (γ) ! look for differences in anisotropy (β)

Subsamples matching $\hat{\langle \gamma \rangle}$ have matching H

Subsamples can have different β_{\perp} & β_{\parallel}

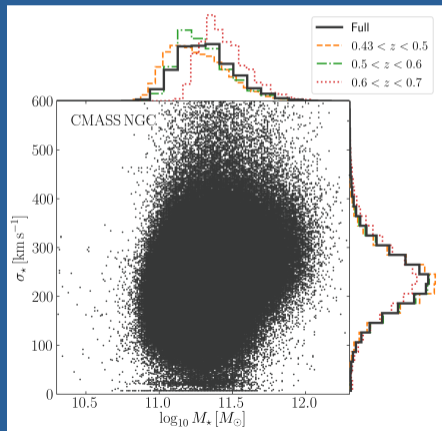
Find subsamples matching d_0 & $\hat{\langle \gamma \rangle}$!

Mismatch d_2 ! evidence $\beta_{\perp} \neq 0$

Galaxy Properties from Portsmouth Group

- velocity dispersion σ_{\star} (1D)
- stellar mass $[M_{\star}]$

Split on γ = split on orientation & galaxy bias $\beta_{\perp}([\gamma])$! different d_0 & d_2



How do we look for AB?

Main idea – split on orientation (γ) ! look for differences in anisotropy (β)

Subsamples matching β have matching H

Subsamples can have different β & β

Find subsamples matching d_0 & β !

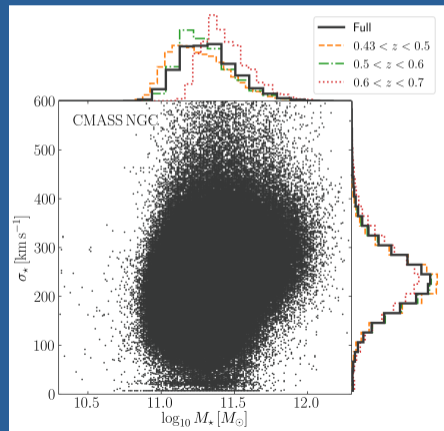
Mismatch d_2 ! evidence $\beta \notin 0$

Galaxy Properties from Portsmouth Group

- velocity dispersion γ (1D)
- stellar mass $[\gamma$

Split on γ = split on orientation & galaxy bias β ($[\gamma$) ! different d_0 & d_2

Use $[\gamma$ to remove mass (β) dependence



How do we look for AB?

Main idea – split on orientation (β) ! look for differences in anisotropy (β_4)

Subsamples matching $\beta^{(\cdot)}$ have matching H

Subsamples can have different β_L & β

Find subsamples matching d_0 & $\beta^{(\cdot)}$!

Mismatch d_2 ! evidence $\beta \notin 0$

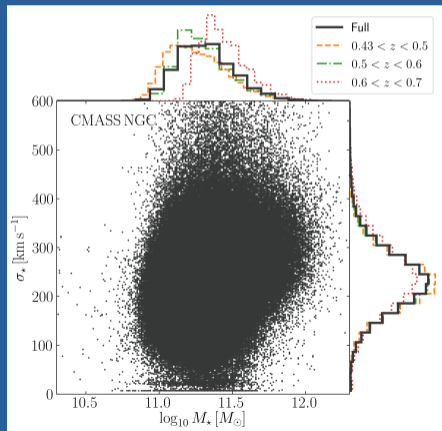
Galaxy Properties from Portsmouth Group

- velocity dispersion β (1D)
- stellar mass $[\beta$

Split on $\beta =$ split on orientation & galaxy bias $\beta_L([\beta)$! different d_0 & d_2

Make subsamples with either

- high $[\beta$, low β bq
- low $[\beta$, high β



How do we look for AB?

Main idea – split on orientation (θ) ! look for differences in anisotropy (d_2)

Subsamples matching θ have matching H

Subsamples can have different d_1 & d_2

Find subsamples matching d_0 & θ !

Mismatch d_2 ! evidence $d_2 \neq 0$

Galaxy Properties from Portsmouth Group

- velocity dispersion σ (1D)
- stellar mass $[M_\odot]$

Split on θ = split on orientation & galaxy bias $d_1(\theta)$! different d_0 & d_2

Make subsamples with either

- high $[M_\odot]$, low σ bq
- low $[M_\odot]$, high σ

How do we look for AB?

Main idea – split on orientation (θ) ! look for differences in anisotropy (d_2)

Subsamples matching θ have matching H

Subsamples can have different d_1 & d_2

Find subsamples matching d_0 & θ !

Mismatch d_2 ! evidence $d_2 \neq 0$

Ob...@b..C\ -zP θ m

- Need to account for θ -evolution
- Work with percentiles
- Compute percentiles in 30 θ -bins
- Split on percentiles in each θ -bin
- E) matching θ

How do we look for AB?

Main idea – split on orientation (θ) ! look for differences in anisotropy (d_1)

Subsamples matching θ have matching H

Subsamples can have different d_1 & d_2

Find subsamples matching d_0 & θ !

Mismatch d_2 ! evidence $d_1 \neq 0$

Ob...@b ..C\ -zP \ b^bebYsm

How do we look for AB?

Main idea – split on orientation (θ) ! look for differences in anisotropy (d_1)

Subsamples matching $\hat{d}(\theta)$ have matching H

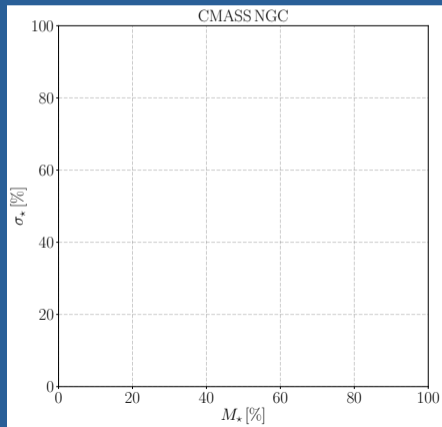
Subsamples can have different d_1 & d_2

Find subsamples matching d_0 & $\hat{d}(\theta)$!

Mismatch d_2 ! evidence $d_1 \neq 0$

Ob...@b..C\ -zP\ b^bebYsm

– Grid of 25 ($\theta; \theta'$) subsamples



How do we look for AB?

Main idea – split on orientation (θ) ! look for differences in anisotropy (ϵ)

Subsamples matching $\hat{(\epsilon)}$ have matching H

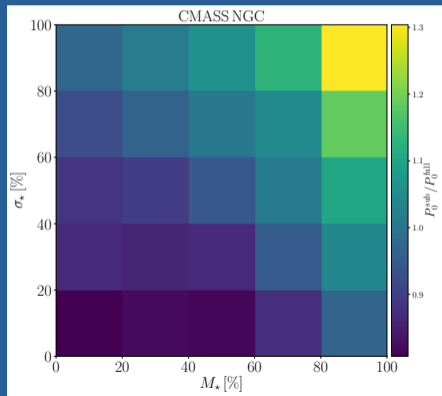
Subsamples can have different ϵ_L & ϵ_T

Find subsamples matching d_0 & $\hat{(\epsilon)}$!

Mismatch d_2 ! evidence $\epsilon \notin 0$

Ob...@b..C\ -z<P\ b^bebYsm

- Grid of 25 ($\theta; [\theta]$) subsamples
- Measure mean amplitude



> Qr /Q r2 HQQF 7Q` "\

J BM B/2 bTHBi Q(M?)Q`Bz2`2M+2b#BM MBbQi`QT

am#b KTH2b K(M?)Q`Bz2`2M+2b#BM ;K i+7BM;

am#b KTH2b + M ? #p2 #Bz2`2Mi

6BM/ bm#b KTH2b K(M?)Q`Bz2`2M+2b#BM;

JBbK i? 2pB/2M#620

> Qr /Q r2 K i+? KQMQTQH2b\

:`B/ Q7 k;8J?) bm#b KTH2b

J2 bm`2 K2 M KTHBim/2

GQ(r?; J?) HQr KTHBim/2

>B;(??; J?) ?B;? KTHBim/2

> Qr /Q r2 HQQF 7Q` "\

J BM B/2 bTHBi Q(M?)Q`Bz2`2M+2b#BM MBbQi`QT

am#b KTH2b K(M?)Q`Bz2`2M+2b#BM; K i+7BM;

am#b KTH2b + M ? #2 #Bz2`2Mi

6BM/ bm#b KTH2b K(M?)Q`Bz2`2M+2b#BM;

JBbK i? 2pB/2M#620

> Qr /Q r2 K i+? KQMQTQH2b\

:`B/ Q7 k;8J?) bm#b KTH2b

J2 bm`2 K2 M KTHBim/2

"mir2 r Mi K i+?BM; KTHBim/25

> Qr /Q r2 HQQF 7Q` "\

J BM B/2 bTHBi Q(M?)Q`Bz2`2M+2b#BM MBbQi`QT

am#b KTH2b K(M)+?? BM;K i+7BM;

am#b KTH2b + M ? #2 #Bz2`2Mi

6BM/ bm#b KTH2b K(M)*3BM;

JBbK i\$? 2pB/2M#620

> Qr /Q r2 K i+? KQMQTQH2b\

6BM HHv b2H2+ib KTH2b rBi?,

?B;J? - HQr

HQJ? - ?B;??

> Qr /Q r2 HQQF 7Q` "\

J BM B/2 bTHBi Q(M?)Q` B2QMF7QM /Bz2`2M+2b#BM MBbQi`QT

am#b KTH2b K(?) +?? B? ;K i+7BM;

am#b KTH2b + M ? #2 #Bz2`2Mi

6BM/ bm#b KTH2\$ K M* 3BM;

JBbK i\$? 2pB/2M#620

> Qr /Q r2 K i+? KQMQTQH2b\

6BM HHv b2H2+ib KTH2b rBi?,

?B;J? - HQr

HQJ? - ?B;??

> Qr /Q r2 HQQF 7Q` "\

J B M B/2 b T H B i Q (M?) Q` B H 2 Q Q F B 7 Q M / B z 2` 2 M + 2 b # B M M B b Q i` Q T

a m # b K T H 2 b K (x) + ?? B M ; K i + 7 B M ;

a m # b K T H 2 b + M ? # 2 # B z 2` 2 M i

6 B M / b m # b K T H 2 b K M * 3 B M ;

J B b K i \$? 2 p B / 2 M # + 2 0

q ? i # Q m i [m / ` m T Q H 2 b \

> Qr /Q r2 HQQF 7Q` "\

J B M B / 2 b T H B i Q (M ?) Q ` B H 2 Q Q F B 7 Q M / B z 2 ` 2 M + 2 b # B M M B b Q i ` Q T

a m # b K T H 2 b K (M) + ? ? B M ; K i + 7 B M ;

a m # b K T H 2 b + M ? # 2 # B z 2 ` 2 M i

6 B M / b m # b K T H 2 B K M * 3 B M ;

J B b K i S ? 2 p B / 2 M # + 2 0

q ? i # Q m i [m / ` m T Q H 2 b \

> Qr /Q r2 HQQF 7Q` "\

J B M B/2 b T H B i Q (M?) Q` B H 2 Q Q F B 7 Q M / B z 2` 2 M + 2 b # B M M B b Q i` Q T

a m # b K T H 2 b K (x) + ?? B M ; K i + 7 B M ;

a m # b K T H 2 b + M ? # 2 # B z 2` 2 M i

6 B M / b m # b K T H 2 B K M * 3 B M ;

J B b K i S ? 2 p B / 2 M # + 2 0

J i + S M (x) ! K i + S

> Qr /Q r2 HQQF 7Q` "\

J B M B / 2 b T H B i Q (M ?) Q ` B H 2 Q M Q F B 7 Q M / B z 2 ` 2 M + 2 b # B M M B b Q i ` Q T

a m # b K T H 2 b K (x) + ? ? B M ; K i + 7 B M ;

a m # b K T H 2 b + M ? # 2 # B z 2 ` 2 M i

6 B M / b m # b K T H 2 B K M * 3 B M ;

J B b K i S ? 2 p B / 2 M # 6 2 0

J i + S M (x) ! K i + S

J B b K i S ? 2 p B / 2 M # 6 2 0

_2bmHib *JaaL:*

PYkyky

_2bmHib GPqw L:*

PYkyky

. 2i2 + iBQM bB; MB} + M + 2

lb2 KQ+F ; H tv + i HQ;b

aTHBi 2 +? KQ+F BM irQ ` M/QK

bm#b KTH2b

*`Qbb@+Q``2H i2 2 +? bm#b KTH2 rBi?

7mHH KQ+F

JBMBK ~~S~~ = 2 S^{bm#} , S^{bm#}

J i+?BM; KQMQTQH2b

rBi?BMi HH b+ H2b

.Bz2`2Mi [m /`m₂ ~~6~~QH2b

K Mvöb r v5

4) #60 #2ir22M bm#b KTH2b

. 2i2+iBQM bB;MB}+ M+2

lb2 KQ+F ; H tv + i HQ;b

aTHBi 2 +? KQ+F BM irQ ` M/QK

bm#b KTH2b

*`Qbb@+Q``2H i2 2 +? bm#b KTH2 rBi?

7mHH KQ+F

JBMBK ~~S~~x=2 S^{bm#} , S^{bm#}

J i+?BM; KQMQTQH2b

rBi?BMi HH b+ H2b

.Bz2`2Mi [m /`m₂QH2b

K Mvöb r v5

4) #60 #2ir22M bm#b KTH2b

* Q K # B M 2 / / 2 i 2 + i B Q M b B ; M B } +

5 m b B M i 0:15 ? J T +¹

P Y k y k y

. B b + m b b B Q M

q2 T` 2 b 2 M i b B ; M B } + M i 2 p B / 2 M + 2 Q 7 " E

. B b + m b b B Q M

q2 T`2b2Mi bB;MB}+ Mi 2pB/2M+2 Q7 " E
q2 b?Qr +Hmbi2`BM; Q7 ; H tB2b /2T2M/
Qi?2` i? M ? HQ K b b

. B b + m b b B Q M

q2 T`2b2Mi bB;MB}+ Mi 2pB/2M+2 Q7 " E

q2 b?Qr +Hmbi2`BM; Q7 ; H tB2b /2T2M/

Qi?2`i? M ? HQ K bb

Pm` /2i2+iBQM Bb i?2 }`bi iQ52t+22/ i?2 M

. B b + m b b B Q M

q2 T`2b2Mi bB;MB}+ Mi 2pB/2M+2 Q7 " E

q2 b?Qr +Hmbi2`BM; Q7 ; H tB2b /2T2M/

Qi?2`i? M ? HQ K bb

Pm` /2i2+iBQM Bb i?2 }`bi iQ52t+22/ i?2 M

q2 QMHv K2 #Q7 2m#b KTH2b- MQi#i? 2 7

. B b + m b b B Q M

q2 T`2b2Mi bB;MB}+ Mi 2pB/2M+2 Q7 " E

q2 b?Qr +Hmbi2`BM; Q7 ; H tB2b /2T2M/

Qi?2`i? M ? HQ K bb

Pm` /2i2+iBQM Bb i?2 }`bi iQ52t+22/ i?2 M

q2 QMHv K2 #b@72bm#b KTH2b- MQi#i5? 2 7

6Q` "Paa ; H tB2b r2 }M/ bm#b 0K1T H02b rB

. B b + m b b B Q M

q2 T`2b2Mi bB;MB}+ Mi 2pB/2M+2 Q7 " E
q2 b?Qr +Hmbi2`BM; Q7 ; H tB2b /2T2M/
Qi?2`i? M ? HQ K bb
Pm` /2i2+iBQM Bb i?2 }`bi iQ52t+22/ i?2 M
q2 QMHv K2 #bQ7 2m#b KTH2b- MQi#i5? 2 7
6Q` "Paa ; H tB2b r2 }M/ bm#b 0K1T 02b rB
6Q` ? HQb BM bBKb r2 7QmM# / b1m #2b KTH

. B b + m b b B Q M

q2 T`2b2Mi bB;MB}+ Mi 2pB/2M+2 Q7 " E
q2 b?Qr +Hmbi2`BM; Q7 ; H tB2b /2T2M/
Qi?2`i? M ? HQ K bb
Pm` /2i2+iBQM Bb i?2 }`bi iQ52t+22/ i?2 M
q2 QMHv K2 #bQ7 2m#b KTH2b- MQi#i5? 2 7
6Q` "Paa ; H tB2b r2 }M/ bm#b 0K1T 02b rB
6Q` ? HQb BM bBKb r2 7QmM# / b1m #2b KTH
JBb HB;MK2Mi Q7 ; H tB2b M/ ? HQb /2

* Q M b 2 [m 2 M + 2 b 7 Q ` . 1 a A

J B M [m 2 b i B Q M b , A b i ` ; 2 i b 2 H 2 + i B Q M Q 7 s Q ` B 2 M i i B Q M / 2 T 2

. 1 a A 1 K B b b B Q M G B M 2 : H t B 2 b U 1 G : b V

* Q M b 2 [m 2 M + 2 b 7 Q ` . 1 a A

J B M [m 2 b i B Q M b , A b i ` ; 2 i b 2 H 2 + i B Q M Q 7 s Q ` B 2 M i i B Q M / 2 T 2

. 1 a A 1 K B b b B Q M G B M 2 : H t B 2 b U 1 G : b V
6 B # 2 ` K ; M B i m / 2 p b X K Q / 2 H K ; M B i m / 2 b

* $QM b^2 [m^2 M + 2b^7 Q] \cdot 1 a A$

J $BM [m^2 b i B Q M b, A b i^2; 2i b^2 H^2 + i B Q M Q^7 s Q^2 B^2 M i i B Q M / 2 T^2$

. $1 a A 1 K B b b B Q M G B M^2 : H t B^2 b U 1 G : b V$

$6 B \#^2 K ; M B i m / 2 p b X K Q / 2 H K ; M B i m / 2 b$

$6 B M i ; H t B^2 b M^2 i^2 / 2 i^2 + i B Q M i^2 b^2 Q H / + Q m H /$

* $QM b^2 [m^2 M + 2b^7 Q] \cdot 1 a A$

J $BM [m^2 b i B Q M b, A b i^2; 2i b^2 H^2 + i B Q M Q^7 s Q^2 B^2 M i i B Q M / 2 T^2$

. $1 a A 1 K B b b B Q M G B M^2 : H t B^2 b U 1 G : b V$

$6 B \# 2^2 K ; M B i m / 2 p b X K Q / 2 H K ; M B i m / 2 b$

$6 B M i ; H t B^2 b M^2 i^2 / 2 i^2 + i B Q M i^2 b^2 Q H / + Q m H /$

$6 B M i ; H t B^2 b^2 H b Q K Q^2 M m K^2 Q m b^5$

* Q M b 2 [m 2 M + 2 b 7 Q ` . 1 a A

J B M [m 2 b i B Q M b , A b i ` ; 2 i b 2 H 2 + i B Q M Q 7 s Q ` B 2 M i i B Q M / 2 T 2

. 1 a A 1 K B b b B Q M G B M 2 : H t B 2 b U 1 G : b V

6 B # 2 ` K ; M B i m / 2 p b X K Q / 2 H K ; M B i m / 2 b

6 B M i ; H t B 2 b M 2 ` i ? 2 / 2 i 2 + i B Q M i ? ` 2 b ? Q H / + Q m H / ;

6 B M i ; H t B 2 b ` 2 H b Q K Q ` 2 M m K 2 ` Q m b 5

1 G : b + Q M i B M / k Q i 2 7 + 2 @ Q M i ? M 2 / ; 2 Q M ; H t B 2 b

* Q M b 2 [m 2 M + 2 b 7 Q ` . 1 a A

J B M [m 2 b i B Q M b , A b i ` ; 2 i b 2 H 2 + i B Q M Q 7 s Q ` B 2 M i i B Q M / 2 T 2

. 1 a A 1 K B b b B Q M G B M 2 : H t B 2 b U 1 G : b V

6 B # 2 ` K ; M B i m / 2 p b X K Q / 2 H K ; M B i m / 2 b

6 B M i ; H t B 2 b M 2 ` i ? 2 / 2 i 2 + i B Q M i ? ` 2 b ? Q H / + Q m H / :

6 B M i ; H t B 2 b ` 2 H b Q K Q ` 2 M m K 2 ` Q m b 5

1 G : b + Q M i B M / k Q i 2 7 + 2 @ Q M i ? M 2 / ; 2 Q M ; H t B 2 b

> Q r 2 p 2 ` - 1 G : b H B ; M 2 / r 2 F 2 ` i ? M G _ : b

* Q M b 2 [m 2 M + 2 b 7 Q ` . 1 a A

J B M [m 2 b i B Q M b , A b i ` ; 2 i b 2 H 2 + i B Q M Q 7 s Q ` B 2 M i i B Q M / 2 T 2

. 1 a A 1 K B b b B Q M G B M 2 : H t B 2 b U 1 G : b V

6 B # 2 ` K ; M B i m / 2 p b X K Q / 2 H K ; M B i m / 2 b

6 B M i ; H t B 2 b M 2 ` i ? 2 / 2 i 2 + i B Q M i ? ` 2 b ? Q H / + Q m H / :

6 B M i ; H t B 2 b ` 2 H b Q K Q ` 2 M m K 2 ` Q m b 5

1 G : b + Q M i B M / k Q i 2 7 + 2 @ Q M i ? M 2 / ; 2 Q M ; H t B 2 b

> Q r 2 p 2 ` - 1 G : b H B ; M 2 / r 2 F 2 ` i ? M G _ : b

q 2 ` 2 + m ` ` 2 M i H v H Q Q F B M ; B M i Q 1 G : i ` ; 2 i b i Q b 2 2 ? C

* Q M b 2 [m 2 M + 2 b 7 Q ` . 1 a A

J B M [m 2 b i B Q M b , A b i ` ; 2 i b 2 H 2 + i B Q M Q 7 s Q ` B 2 M i i B Q M / 2 T 2

. 1 a A 1 K B b b B Q M G B M 2 : H t B 2 b U 1 G : b V

6 B # 2 ` K ; M B i m / 2 p b X K Q / 2 H K ; M B i m / 2 b

6 B M i ; H t B 2 b M 2 ` i ? 2 / 2 i 2 + i B Q M i ? ` 2 b ? Q H / + Q m H / :

6 B M i ; H t B 2 b ` 2 H b Q K Q ` 2 M m K 2 ` Q m b 5

1 G : b + Q M i B M / k Q i 2 7 + 2 @ Q M i ? M 2 / ; 2 Q M ; H t B 2 b

> Q r 2 p 2 ` - 1 G : b H B ; M 2 / r 2 F 2 ` i ? M G _ : b

q 2 ` 2 + m ` ` 2 M i H v H Q Q F B M ; B M i Q 1 G : i ` ; 2 i b i Q b 2 2 ? C

. 1 a A G m K B M Q m b _ 2 / : H t B 2 b U G _ : b V

* Q M b 2 [m 2 M + 2 b 7 Q ` . 1 a A

J B M [m 2 b i B Q M b , A b i ` ; 2 i b 2 H 2 + i B Q M Q 7 s Q ` B 2 M i i B Q M / 2 T 2

. 1 a A 1 K B b b B Q M G B M 2 : H t B 2 b U 1 G : b V

6 B # 2 ` K ; M B i m / 2 p b X K Q / 2 H K ; M B i m / 2 b

6 B M i ; H t B 2 b M 2 ` i ? 2 / 2 i 2 + i B Q M i ? ` 2 b ? Q H / + Q m H / :

6 B M i ; H t B 2 b ` 2 H b Q K Q ` 2 M m K 2 ` Q m b 5

1 G : b + Q M i B M / k Q i 2 7 + 2 @ Q M i ? M 2 / ; 2 Q M ; H t B 2 b

> Q r 2 p 2 ` - 1 G : b H B ; M 2 / r 2 F 2 ` i ? M G _ : b

q 2 ` 2 + m ` ` 2 M i H v H Q Q F B M ; B M i Q 1 G : i ` ; 2 i b i Q b 2 2 ? C

. 1 a A G m K B M Q m b _ 2 / : H t B 2 b U G _ : b V

a B K B H ` i Q " P a a - # 2 i i 2 ` K 2 b m ` 2 K 2 M i Q 7 "

* Q M b 2 [m 2 M + 2 b 7 Q ` . 1 a A

J B M [m 2 b i B Q M b , A b i ` ; 2 i b 2 H 2 + i B Q M Q 7 s Q ` B 2 M i i B Q M / 2 T 2

. 1 a A 1 K B b b B Q M G B M 2 : H t B 2 b U 1 G : b V

6 B # 2 ` K ; M B i m / 2 p b X K Q / 2 H K ; M B i m / 2 b

6 B M i ; H t B 2 b M 2 ` i ? 2 / 2 i 2 + i B Q M i ? ` 2 b ? Q H / + Q m H / :

6 B M i ; H t B 2 b ` 2 H b Q K Q ` 2 M m K 2 ` Q m b 5

1 G : b + Q M i B M / K Q i 2 7 + 2 @ Q M i ? M 2 / ; 2 Q M ; H t B 2 b

> Q r 2 p 2 ` - 1 G : b H B ; M 2 / r 2 F 2 ` i ? M G _ : b

q 2 ` 2 + m ` ` 2 M i H v H Q Q F B M ; B M i Q 1 G : i ` ; 2 i b i Q b 2 2 ? C

. 1 a A G m K B M Q m b _ 2 / : H t B 2 b U G _ : b V

a B K B H ` i Q " P a a - # 2 i i 2 ` K 2 b m ` 2 K 2 M i Q 7 "

` 2 G _ : b / m b i 7 ` 2 2 \

$$P_i = 2 \left(\frac{Q M}{b} + \frac{m^2 M}{2b} \right)$$

$$P_i = \frac{Q m T b f + H m b i^2}{b} = \frac{Q m M}{B M} = \frac{2}{b} B i @$$

Other consequences

Kp~esw<YszCp Hb~^@S q@PSHGe <C

- ... in simulations exhibit strong AB signal in their clustering

Other consequences

Kp~eswYszCp Hb~^@S qC@PSHGe <C

– ... in simulations exhibit strong AB signal in their clustering

; b~Y@4C- ^ Ss~CHbqpr? ..SP |c<\ RzC^sS%b - eeSL

Other consequences

K ϕ -esw \leq YszC ϕ Hb \sim ^@S \wedge qC@PSH ϕ e- <C

- ... in simulations exhibit strong AB signal in their clustering

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- ... due to HI self-absorption, provided HI aligned with halos

Other consequences

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iii

Summary

$\int b^Q \langle YqP \cdot Y \text{ edge } CqS - q \langle bqYzC@ \dots SP YqCq \langle YzS@ Y' C@S$
a $q^z z^b @Cei sCz^b C Czs - ^@zS@ Y- YL^ \ Cz bHP- YswL- Y\ddagger S=$
$$\frac{s}{L} = (4L - 4/3 + (H - 4)^2) \setminus$$

„ C” $^@zCq \ s \ YW 4 \ eqsCzS^ sS \sim Yz^s - ^@3arr L- Y\ddagger S$
- re $Yz L- Y\ddagger S 4 \ sC@b^ YCbHsLPz fCb < S% @SeCpsb^ ? - ^@[?$
rS e $YzCsz \ s \sim 4s \setminus eYs \setminus -zPSL d_0 \ E) \setminus -zPSL d_2$

K- $Y\ddagger \% YszCq^L @CeC^@s b^ bzPCq Y \langle YedpeCqS \rangle ^bz U sz P- Y \setminus -ss=$

First detection of galaxy assembly bias to exceed 5 !

Problem for RSD since 4_i completely degenerate with H

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Other splits and approaches

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Kaiser model for multipoles assuming $\mathcal{A} = 0$

AO+2020

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AO+2020

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- Perhaps explaining previous results...

AO+2020

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