Leonardo Senatore (Stanford)

Effective Field Theory of Large Scale Structure

the way to go for inflation

A talk about

- A nice EFT
- Some GR
- high-energy techniques applied to a novel setting

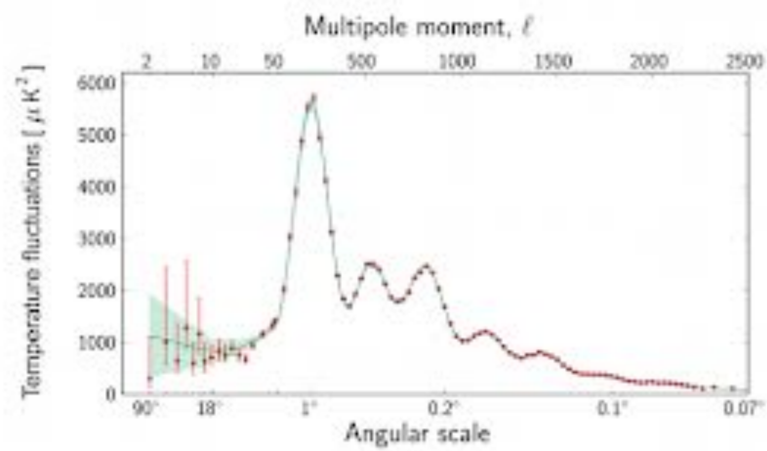
- what the 10 year future of inflationary cosmology stands on
  - as I am now going to argue

How do we probe inflation

- The only observable we are testing from the background solution is

  $\Omega_K \lesssim 3 \times 10^{-3}$

- All the rest, comes from the fluctuations



- For the fluctuations
  - they are primordial
  - they are scale invariant
  - they have a tilt $\quad n_s - 1 \simeq -0.04 \sim \mathcal{O}\left(\frac{1}{N_e}\right)$
  - they are quite gaussian

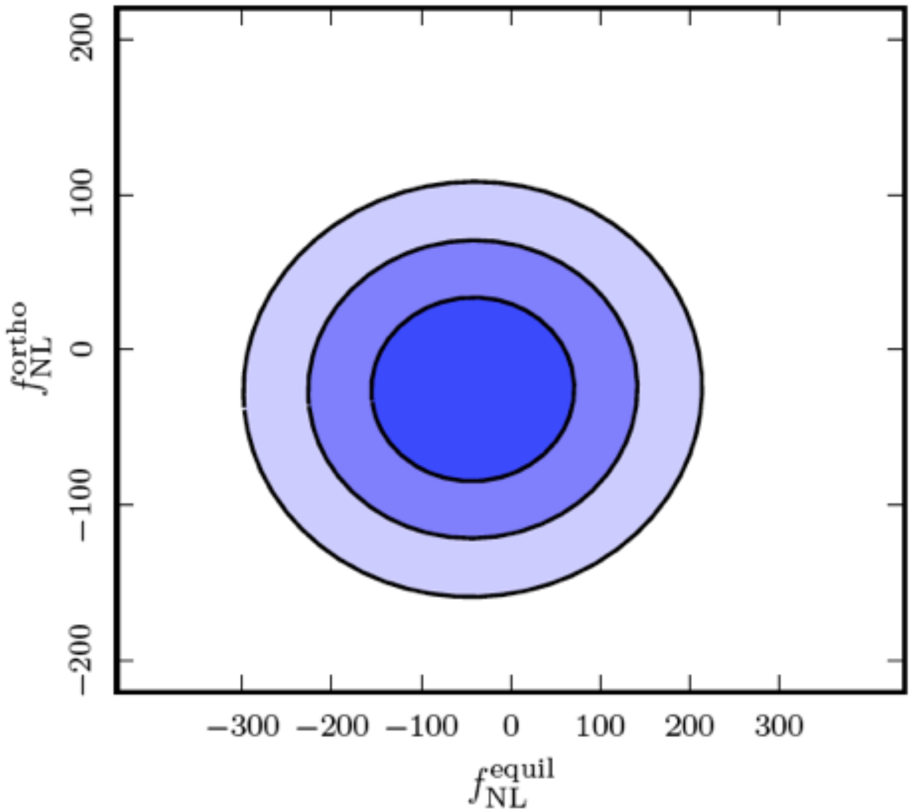
$$\mathrm{NG} \sim \frac{\langle \zeta^3 \rangle}{\langle \zeta^2 \rangle^{3/2}} \lesssim 10^{-3}$$

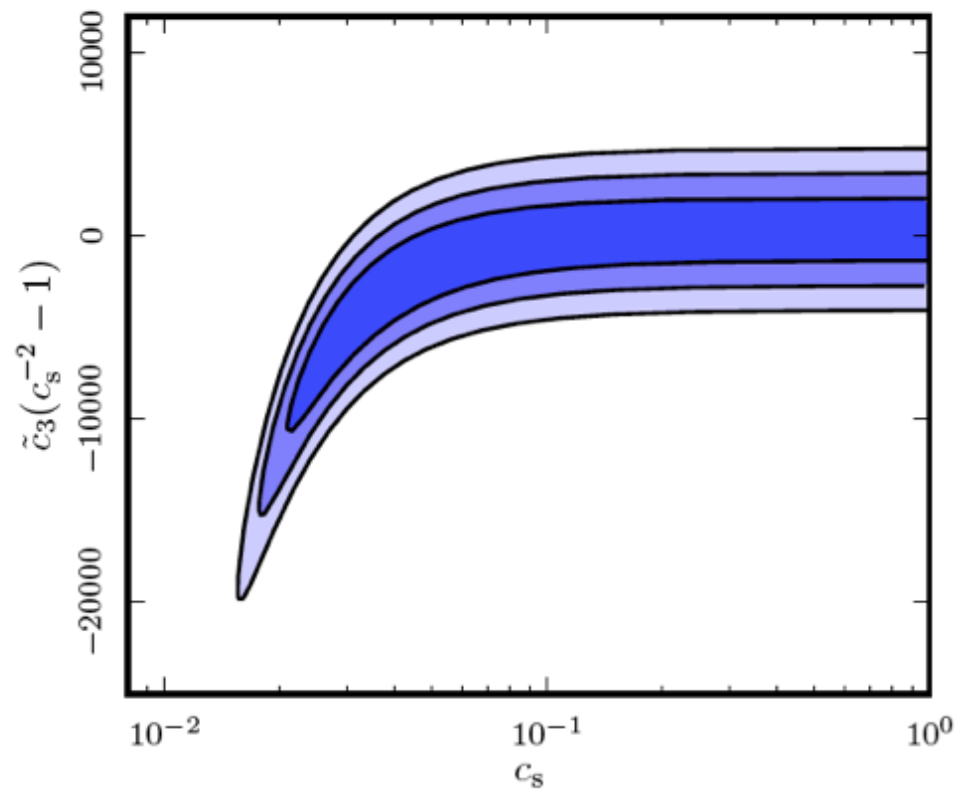
  - both scalar and maybe tensors

Limits in terms of parameters of a Lagrangian

- $$S = \int d^4x \sqrt{-g}\left[-\frac{M_{\rm Pl}^2 \dot H}{c_s^2}\left(\dot\pi^2 - c_s^2\frac{(\partial_i\pi)^2}{a^2}\right) + (M_{\rm Pl}^2\dot H)\frac{1-c_s^2}{c_s^2}\left(\frac{\dot\pi(\partial_i\pi)^2}{a^2} + \frac{A}{c_s^2}\dot\pi^3\right) + \cdots\right]$$

with C. Cheung, P. Creminelli, L. Fitzpatrick, J. Kaplan **JHEP 2008**





with Smith and Zaldarriaga, **JCAP2010**
Planck Collaboration **2013**

- these are limits on the cutoff of the theory $\sim \dfrac{\dot\pi^3}{\Lambda^2}$

What has Planck done to theory?

- Planck improve limits wrt WMAP by a factor of ~3.
- Since $\mathrm{NG} \sim \frac{H^2}{\Lambda_U^2} \quad \Rightarrow \quad \Lambda_U^{\mathrm{min,\,Planck}} \simeq 2\,\Lambda_U^{\mathrm{min,\,WMAP}}$
- Given the absence of known or nearby threshold, this is not much.
- Planck was great
- but Planck was not good enough
  - not Plank's fault, but Nature's faults
    - Please complain with Nature
- Planck was an opportunity for a detection, not much an opportunity to change the theory in absence of detection (luckily WMAP had a tilt a 2.5 σ , so we got to 6 σ)
- On theory side, little changes
  - contrary for example to LHC, which was crossing thresholds
    - Any result from LHC is changing the theory

Cosmology is going to change in a few months

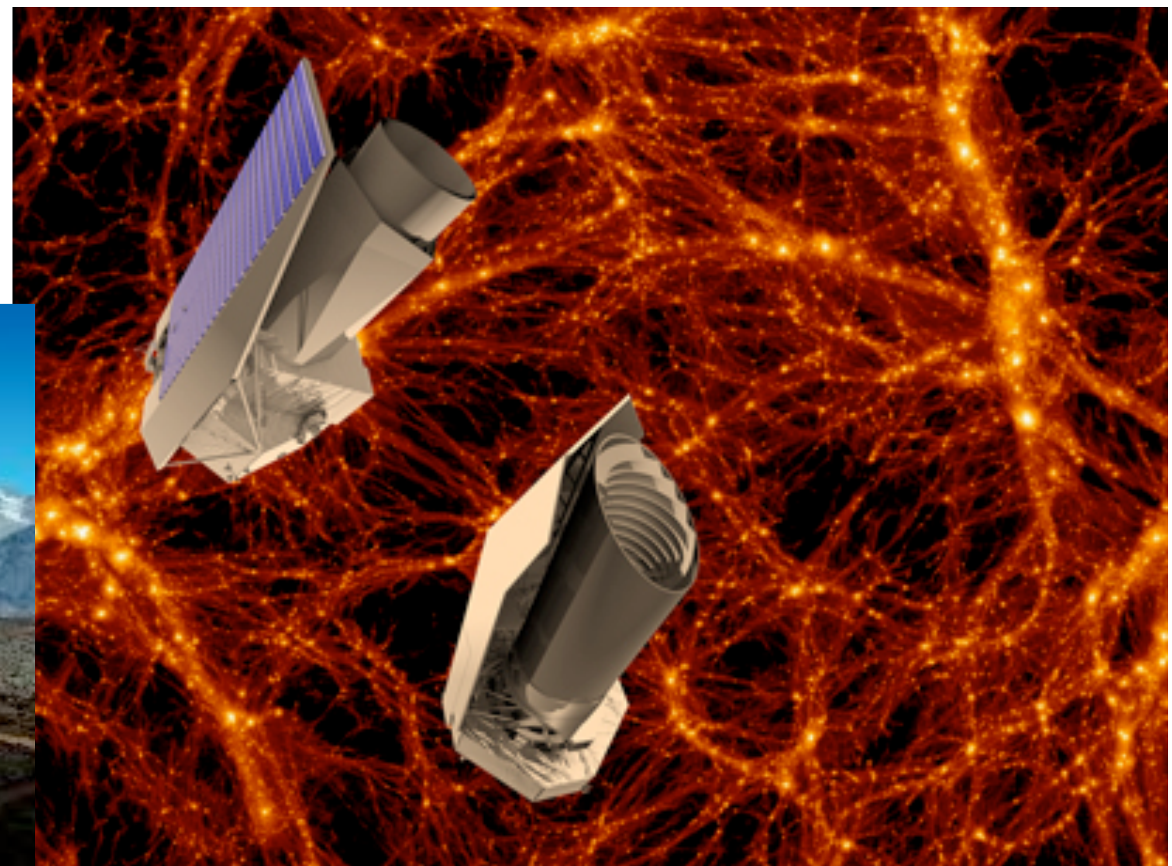
- Tremendous progress has been made through observation of the primordial fluctuations
- In order to increase our knowledge of Inflation, we need more modes
- Planck will soon have observed all the modes from the CMB
- and then what?
- I will assume we are not lucky
  - no B-mode detection
  - no signs from the beginning of inflation
- Unless we find a way to get more modes, the game is over
- Large Scale Structures offer the only medium-term place for hunting for more modes
  - but we are compelled to understand them
    - I do not think, so far, we understand them well enough

What is next?

- Euclid and LSST like: this is our only next chance
  - we need to understand how many modes are available

$$\text{Number of modes} \sim \left(\frac{k_{\max}}{k_{\min}}\right)^3$$

  - Need to understand short distances
  - Similar as from LEP to LHC

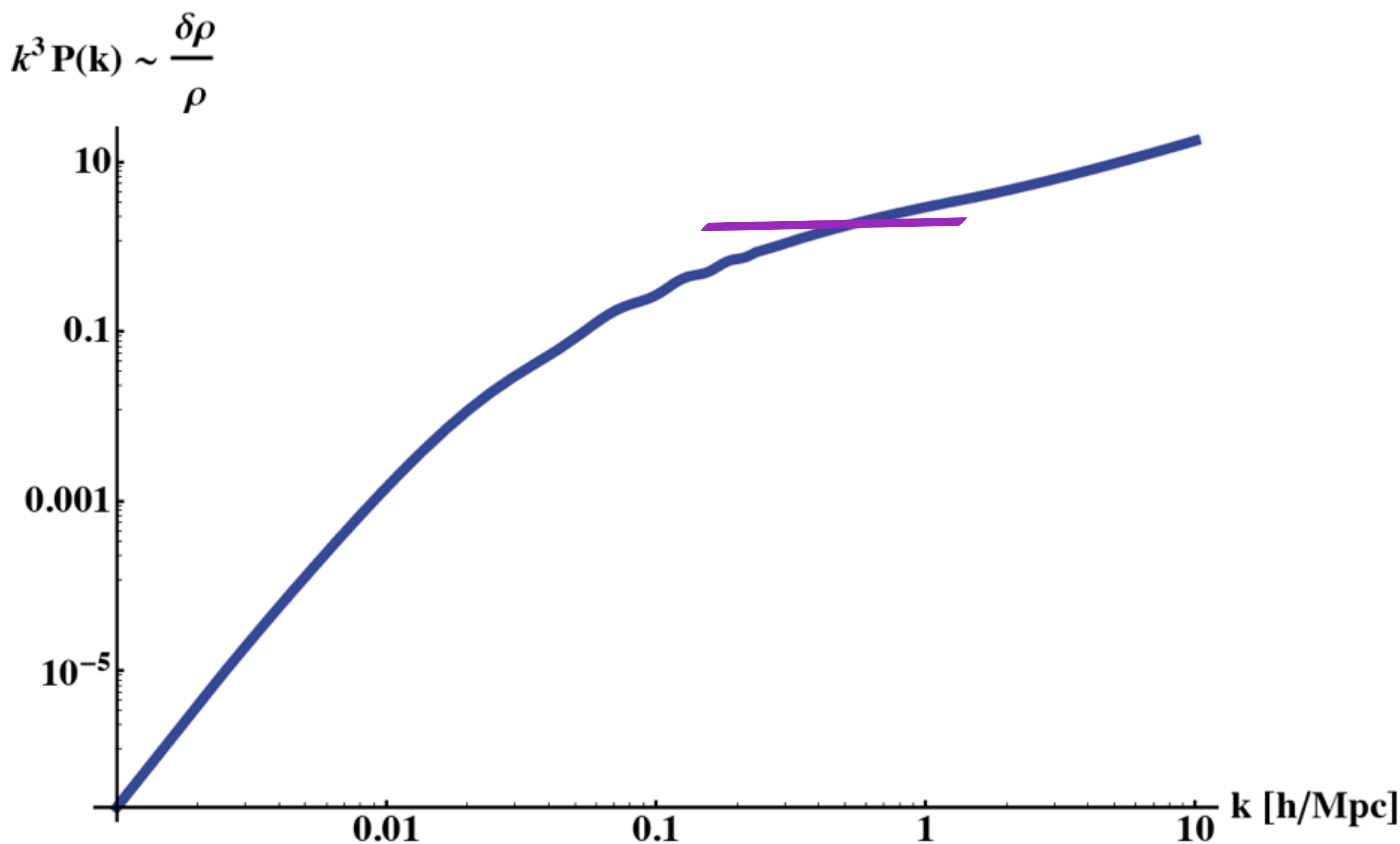


The Effective Field Theory of Cosmological Large Scale Structures

Redshift Space distortions in the EFTofLSS	with Zaldarriaga **1409**
Bias in the EFTofLSS	me alone **1406**
The one-loop bispectrum in the EFTofLSS	with Angulo, Foreman, Schmittful **1406** see also Baldauf, Mirbabayi, Mercolli,Pajer **1406**
The IR-resummed EFTofLSS	with Zaldarriaga **1404**
The Lagrangian-space EFTofLSS	with Porto and Zaldarriaga **JCAP1405**
The EFTofLSS at 2-loops	with Carrasco, Foreman and Green **JCAP1407**
The 2-loop power spectrum and the IR safe integrand	with Carrasco, Foreman and Green **JCAP1407**
The Effective Theory of Large Scale Structure (EFTofLSS)	with Carrasco and Hertzberg **JHEP 2012**
Cosmological Non-linearities as an Effective Fluid	with Baumann, Nicolis and Zaldarriaga **JCAP 2012**

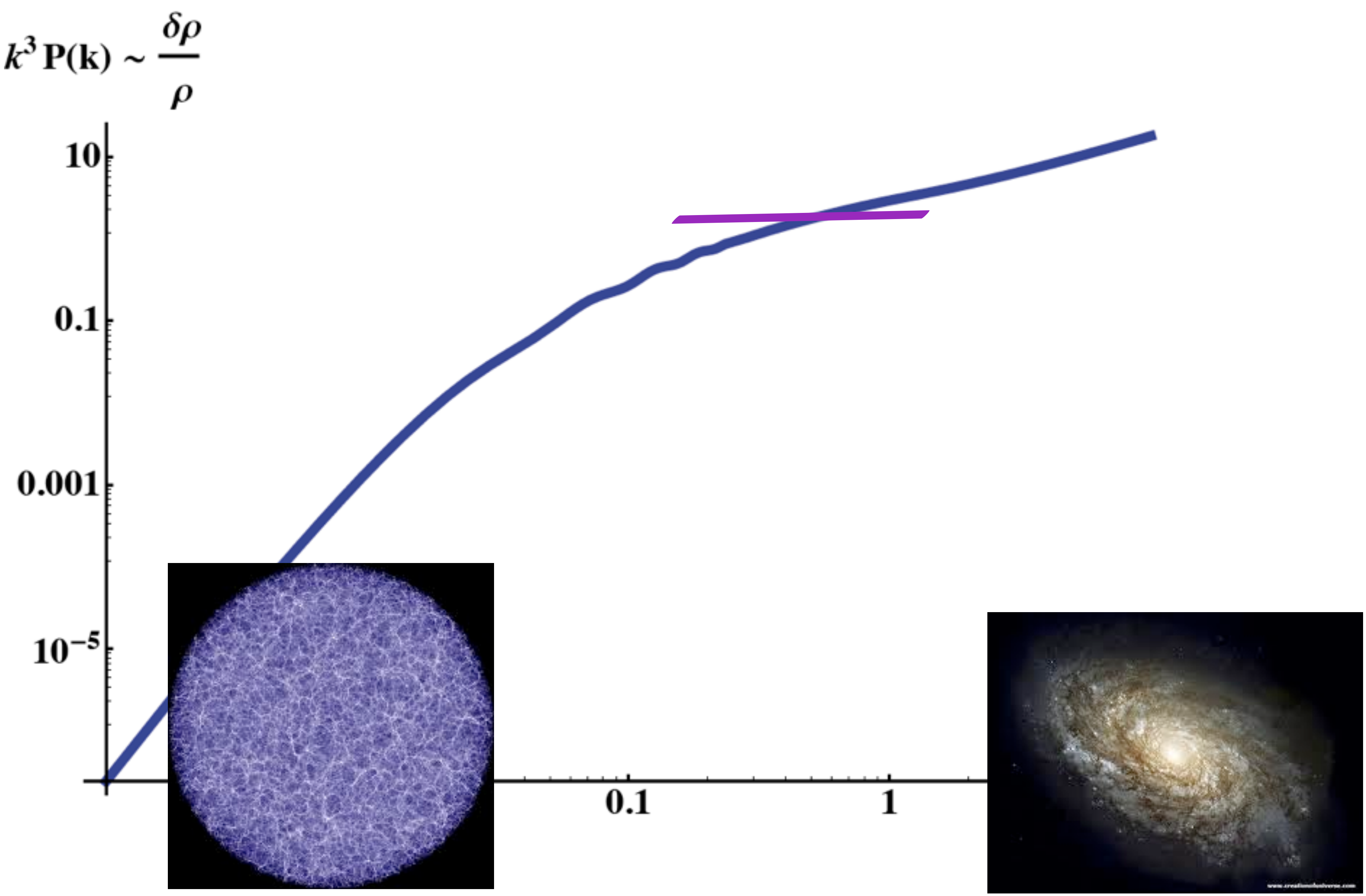
A well defined perturbation theory

- Non-linearities at short scale



A well defined perturbation theory

- Non-linearities at short scale



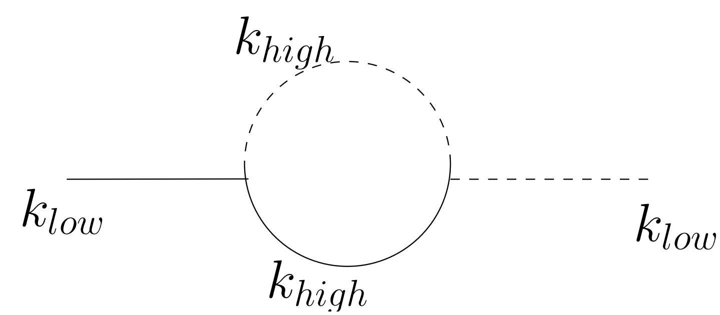
A well defined perturbation theory

- Standard perturbation theory is not well defined
- Standard techniques
  - perfect fluid $\quad \dot{\rho} + \partial_i \left(\rho v^i \right) = 0 \ ,$
  - expand in $\quad \delta \sim \dfrac{\delta \rho}{\rho}$ and solve iteratively

$$\delta^{(n)} \sim \int \text{GreenFunction} \times \text{Source}^{(n)} \left[\delta^{(1)}, \delta^{(2)}, \ldots, \delta^{(n-1)} \right]$$

$$\Rightarrow \quad \langle \delta_k^{(2)} \delta_k^{(2)} \rangle \sim \int d^3k' \ \langle \delta_{k-k'}^{(1)} \delta_{k-k'}^{(1)} \rangle \ \langle \delta_{k'}^{(1)} \delta_{k'}^{(1)} \rangle$$

- Perturbative equations break in the UV
  - $\delta \sim \dfrac{k}{k_{NL}} \gg 1 \quad \text{for} \quad k \gg k_{NL}$
  - no perfect fluid if we truncate

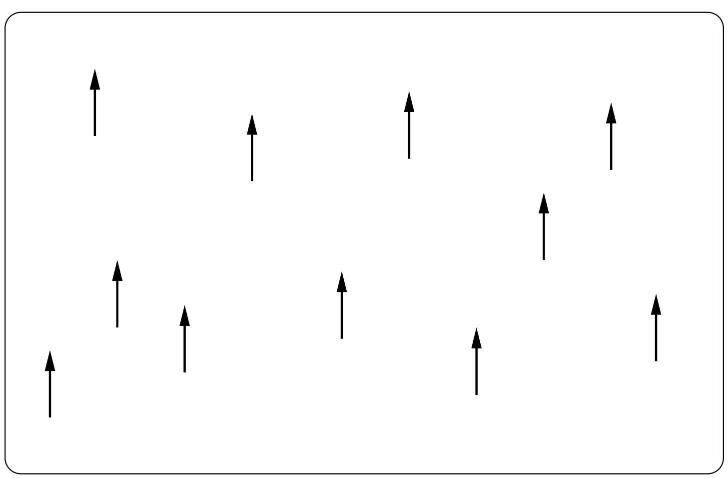


Idea of the Effective Field Theory

Consider a dielectric material

- Very complicated on atomic scales d_{atomic}
- On long distances $d \gg d_{\text{atomic}}$
  - we can describe atoms with their gross characteristics
    - polarizability $\vec{d}_{\text{dipole}} \sim \alpha\, \vec{E}_{\text{electric}}$: average response to electric field
  - we are led to a uniform, smooth material, with just some macroscopic properties
    - we simply solve Maxwell dielectric equations, we do not solve for each atom.
- The universe looks like a dielectric

Dielectric Fluid



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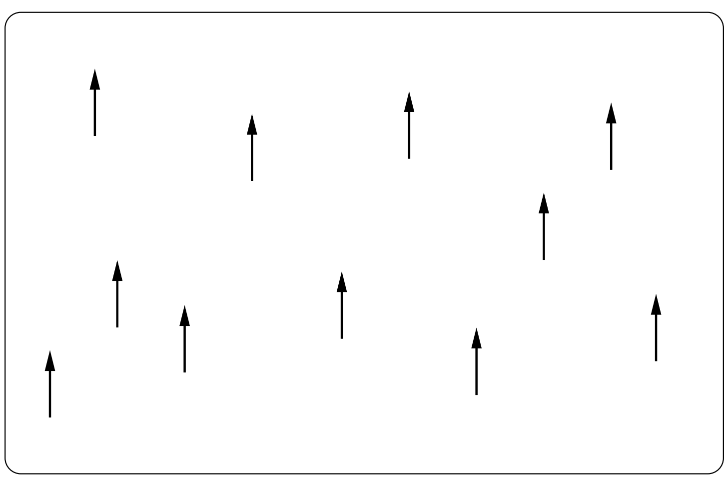
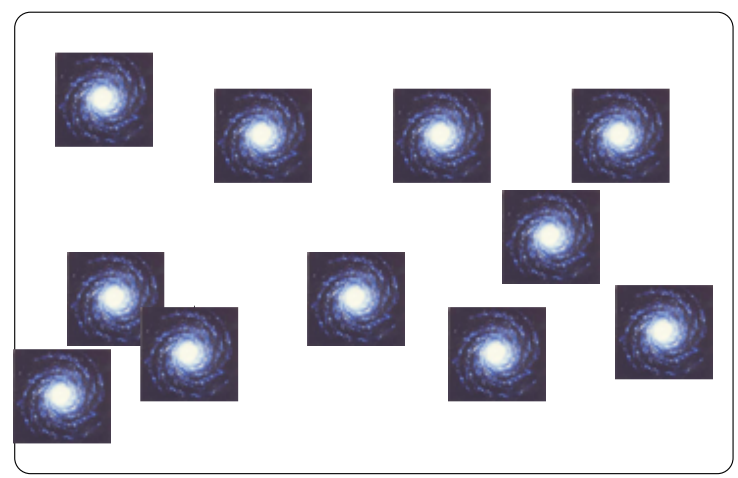
$$\text{EM} \quad \longrightarrow \quad \text{GR}$$

Dielectric Fluid

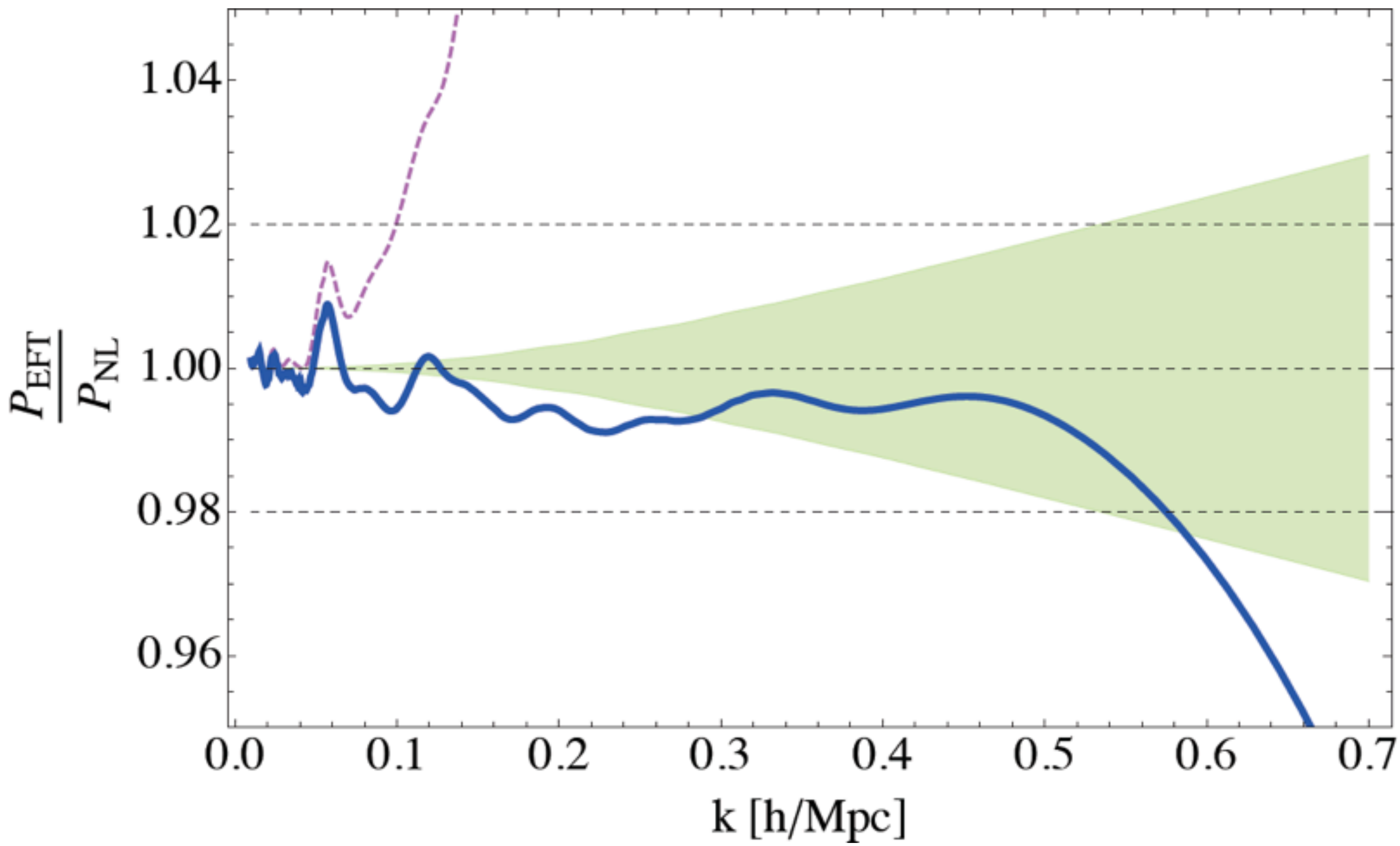


Dielectric Fluid



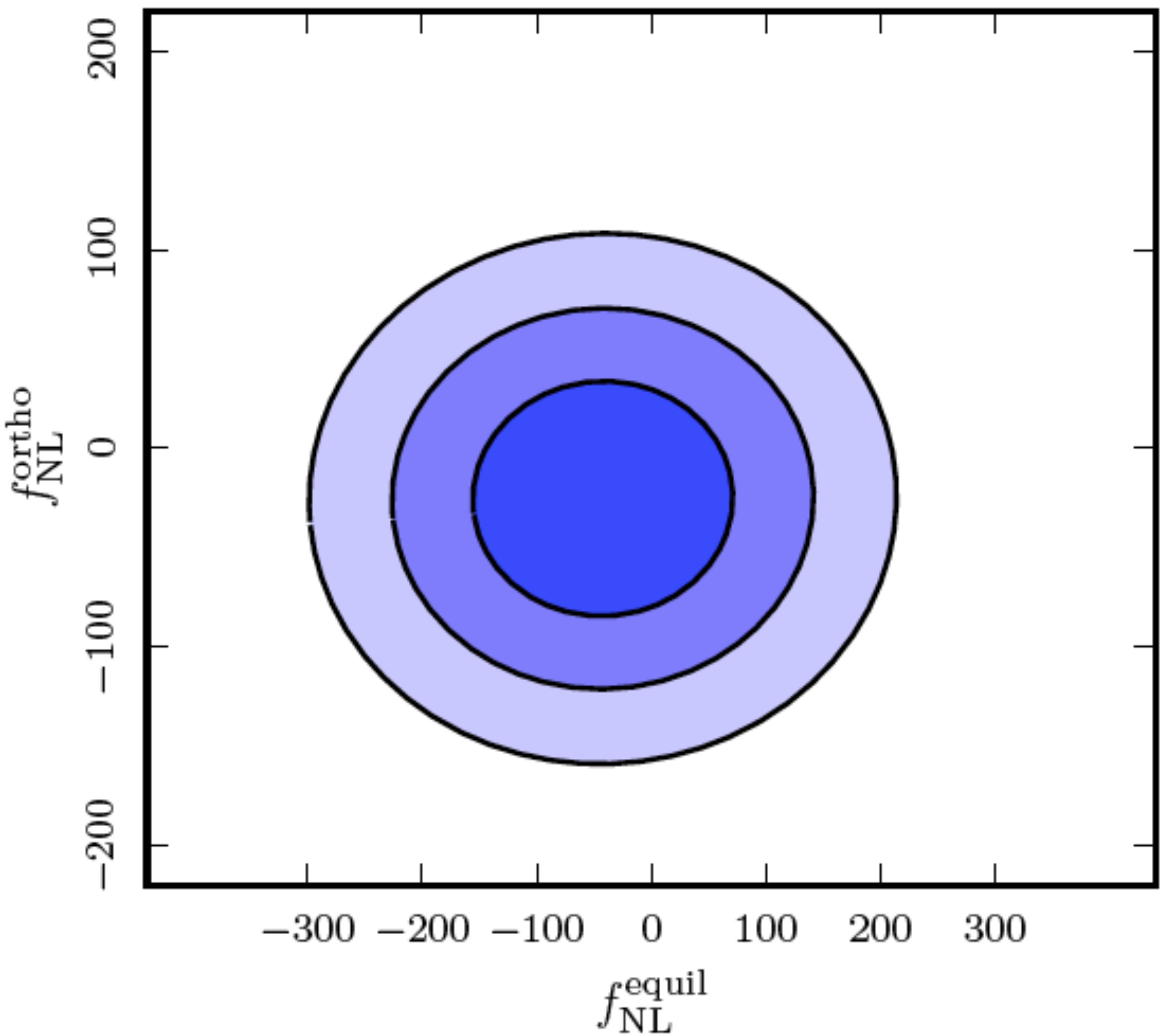
Bottom line result

- A well defined perturbation theory
- 2-loop in the EFT, with IR resummation

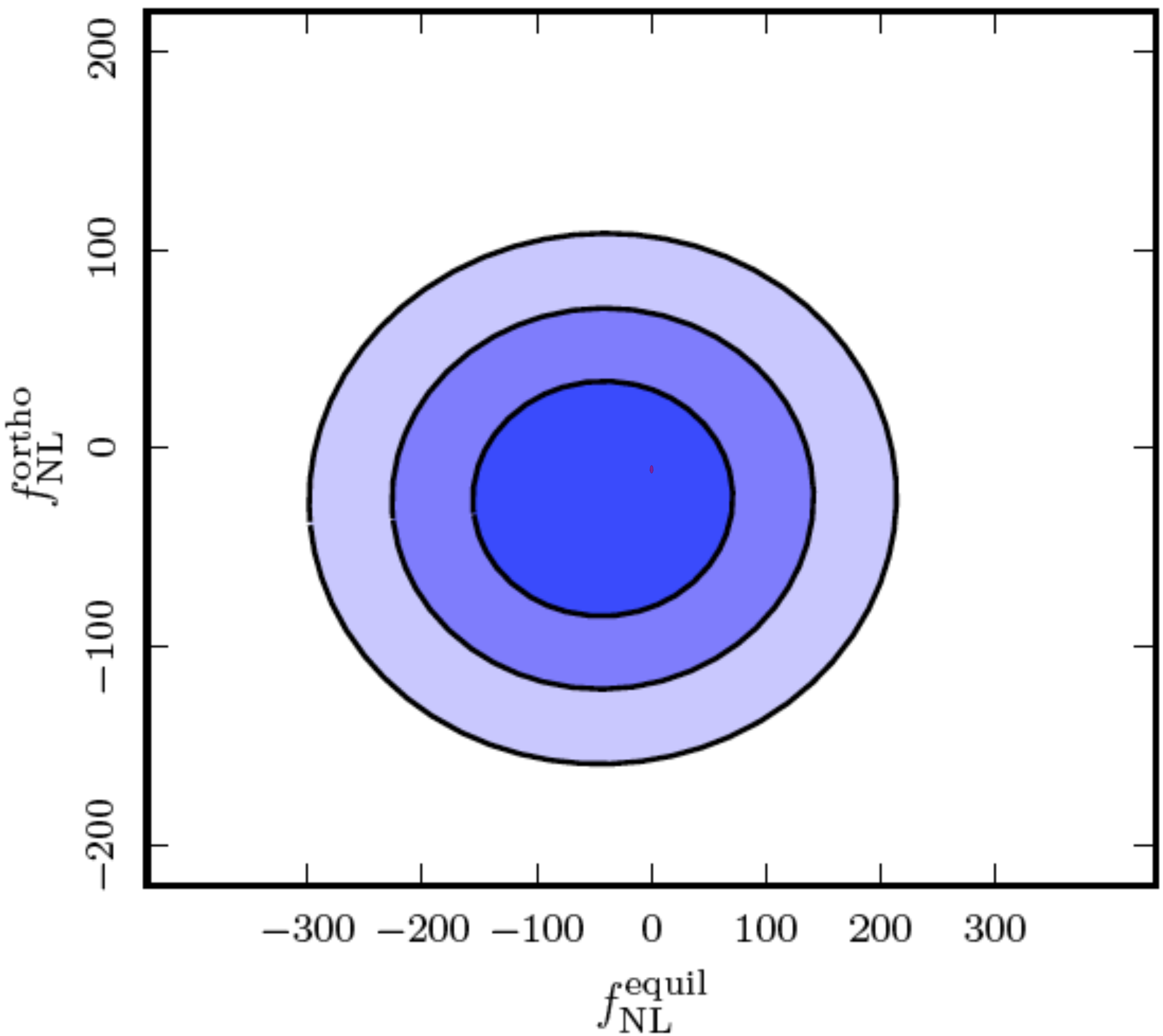


- Data go as $k_{\rm max}^3$: naively factor of 200 more modes than before

With this

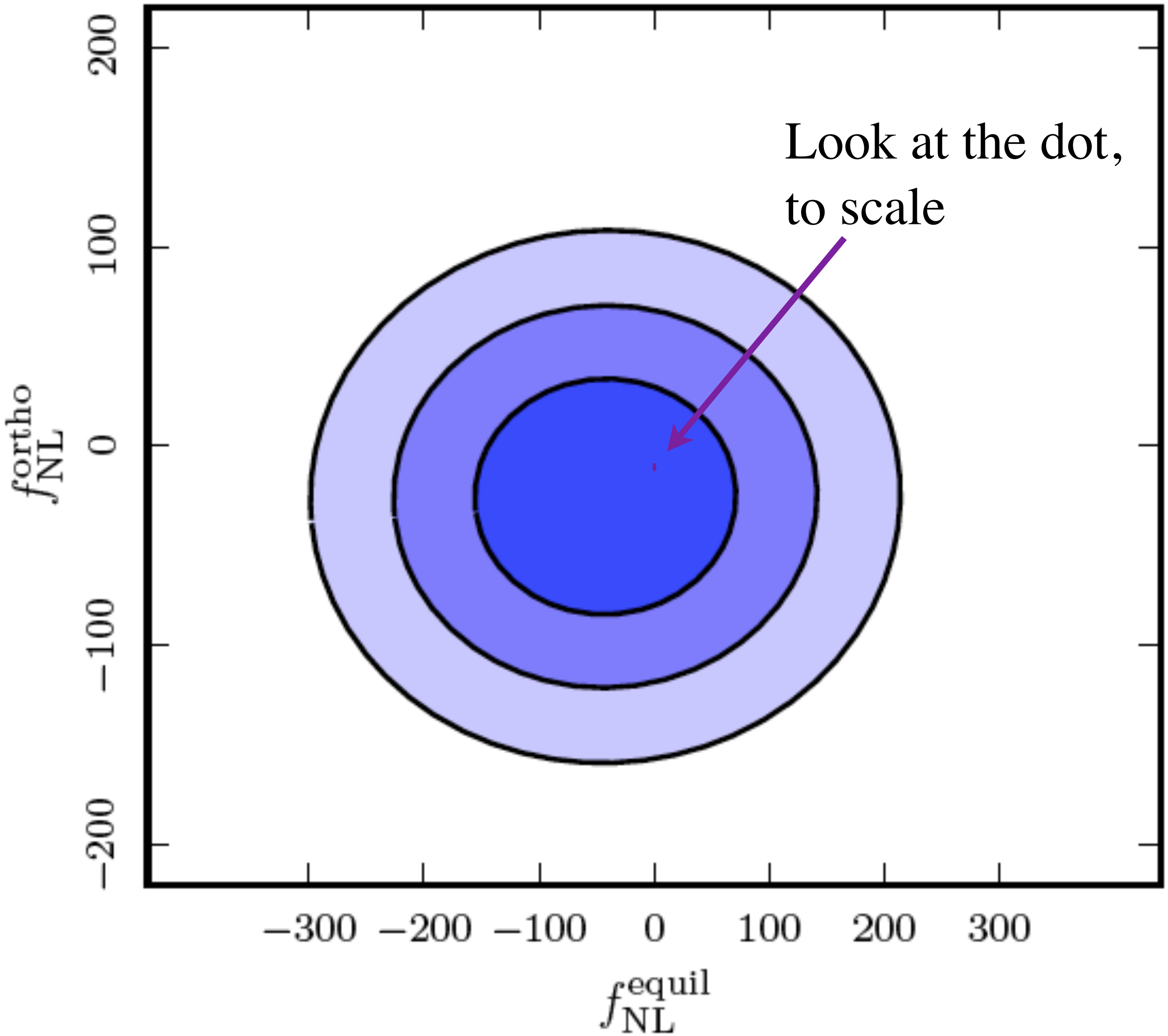


With this



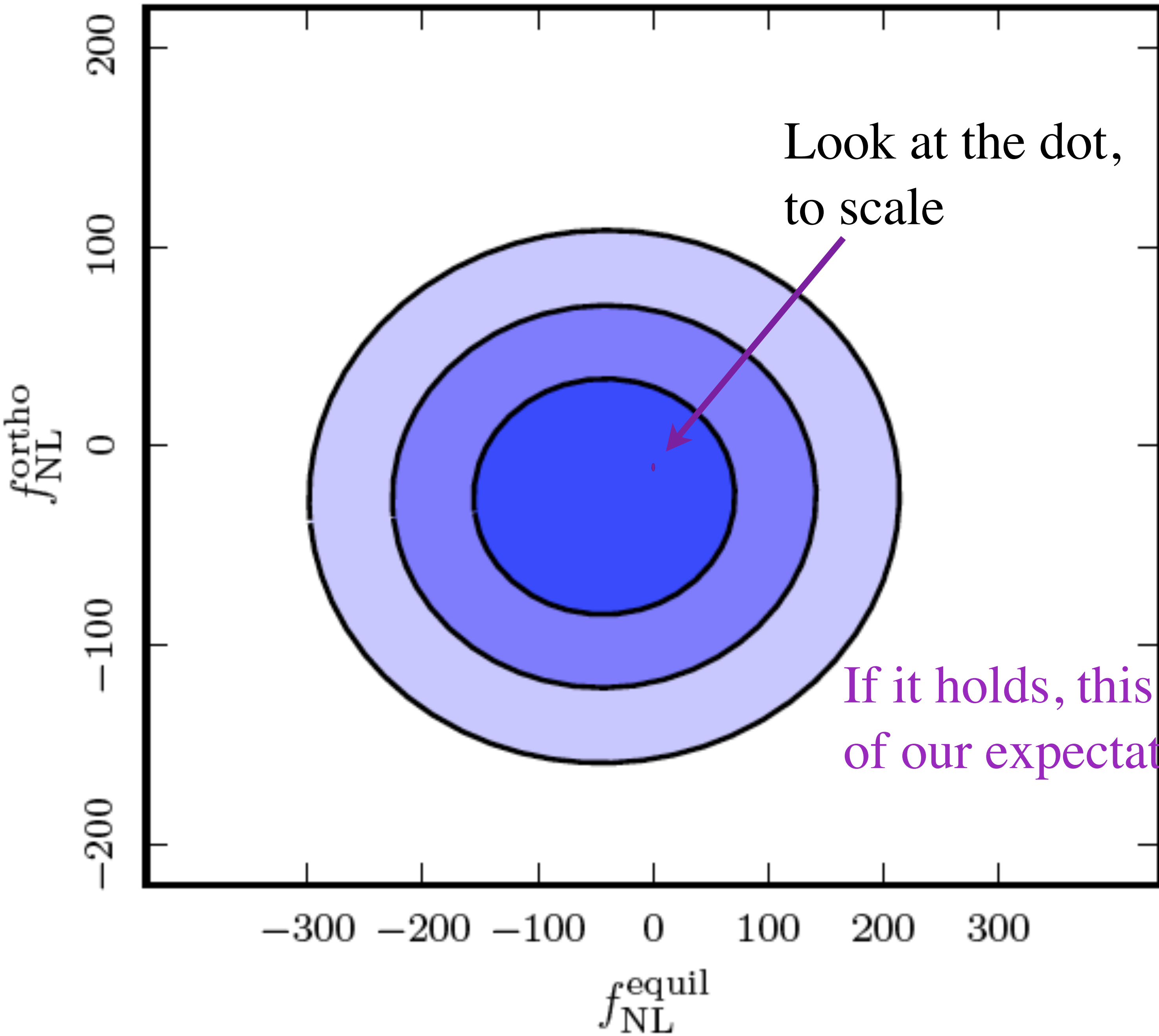
$f_{\mathrm{NL}}^{\mathrm{ortho}}$
200
100
0
−100
−200
−300
−200
−100
0
100
200
300
$f_{\mathrm{NL}}^{\mathrm{equil}}$

With this



Look at the dot,
to scale
200
100
0
−100
−200
f_{NL}^{ortho}
−300
−200
−100
0
100
200
300
f_{NL}^{equil}

With this



If it holds, this is a revolution
of our expectations

Construction of the Effective Field Theory

Point-like Particle versus Extended Objects

- On short distances, we have point-like particles
  - they move

$$\frac{d^2\vec{z}(\vec{q},\eta)}{d\eta^2}+\mathcal{H}\frac{d\vec{z}(\vec{q},\eta)}{d\eta}=-\vec{\partial}_x\Phi[\vec{z}(\vec{q},\eta)]$$

  - induce overdensities

$$1+\delta(\vec{x},\eta)=\int d^3q\;\delta^{(3)}(\vec{x}-\vec{z}(\vec{q},\eta))$$

  - Source gravity

$$\partial^2\Phi(\vec{x})=\mathcal{H}^2\delta(\vec{x})$$

Point-like Particle versus Extended Objects

- But we cannot describe point-like particles: we need to focus on long distances.
  - We deal with Extended objects
    - they move differently:

$$\frac{d^2\vec{z}(\vec{q},\eta)}{d\eta^2}+\mathcal{H}\frac{d\vec{z}(\vec{q},\eta)}{d\eta}=-\vec{\partial}_x\Phi[\vec{z}(\vec{q},\eta)]$$

Point-like Particle versus Extended Objects

- But we cannot describe point-like particles: we need to focus on long distances.
  - We deal with Extended objects
    - they move differently:

$$\frac{d^2\vec{z}_L(\vec{q},\eta)}{d\eta^2}+\mathcal{H}\frac{d\vec{z}_L(\vec{q},\eta)}{d\eta} \quad = \quad -\vec{\partial}_x\left[\Phi_L[\vec{z}_L(\vec{q},\eta)]+\frac{1}{2}Q^{ij}(\vec{q},\eta)\partial_i\partial_j\Phi_L[\vec{z}_L(\vec{q},\eta)]+\cdots\right]+\vec{a}_S(\vec{q},\eta)$$

Point-like Particle versus Extended Objects

- They induce number over-densities and real-space multipole moments

$$
\begin{aligned}
1+\delta_{n,L}(\vec{x},\eta) &\equiv \int d^3\vec{q}\;\delta^3(\vec{x}-\vec{z}_L(\vec{q},\eta))\;, \\
\mathcal{Q}^{i_1\ldots i_p}(\vec{x},\eta) &\equiv \int d^3\vec{q}\;Q^{i_1\ldots i_p}(\vec{q},\eta)\delta^3(\vec{x}-\vec{z}_L(\vec{q},\eta))
\end{aligned}
$$

- they source gravity with the `overall’ mass

$$
\partial_x^2\Phi_L=\frac{3}{2}\mathcal{H}^2\Omega_m\left(\delta_{n,L}(\vec{x},\eta)+\frac{1}{2}\partial_i\partial_j\mathcal{Q}^{ij}(\vec{x},\eta)-\frac{1}{6}\partial_i\partial_j\partial_k\mathcal{Q}^{ijk}(\vec{x},\eta)+\cdots\right)\equiv\frac{3}{2}\mathcal{H}^2\Omega_m\delta_{m,L}(\vec{x},\eta)
$$

$$
\sim \text{Energy}_{\text{electrostatic}} = q\,V+\vec{d}\cdot\vec{E}+\ldots
$$

- These equations can be derived from smoothing the point-particle equations
  - but actually these are the assumption-less equations

How do we treat the new terms?

- Similar to treatment of material polarizability: $\vec{d}_{\text{dipole}} \sim \vec{d}_{\text{intrinsic}} + \alpha\, \vec{E}$
- Take moments:

$$Q^{ij} = \langle Q^{ij} \rangle_S + Q^{ij}_{\mathcal{S}} + Q^{ij}_{\mathcal{R}}$$

- Expectation value $\langle Q^{ij} \rangle_{\mathcal{S}} = l_S^2(\eta)\delta_{ij}$
- Response (non-local in time) $Q_{ij,\mathcal{R}} \sim l_1(\eta)^2\, \partial_i \partial_j \Phi_L(\vec{z}_L(\vec{q}, \eta)$
- Stochastic noise

$$\langle Q_{\mathcal{S}} \rangle = 0 \qquad \langle Q_{\mathcal{S}} Q_{\mathcal{S}} \ldots \rangle \neq 0$$

- Overall

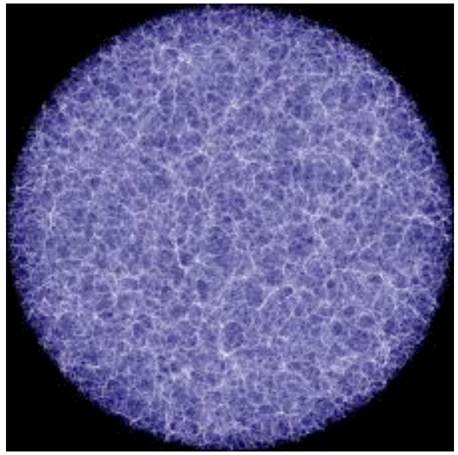
$$Q_{ij}(\vec{x}, t) = l_0^2(t)\, \delta_{ij} + l_1^2(t)\, \partial_i \partial_j \Phi(\vec{x}, t) + \ldots$$

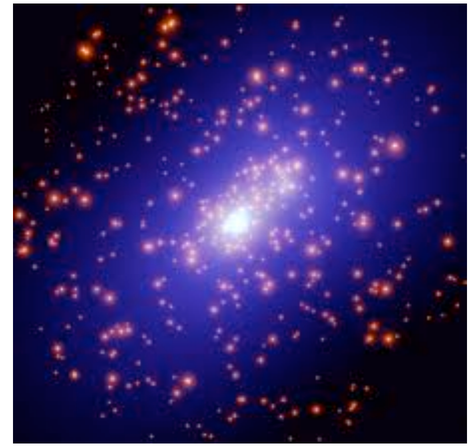
- In summary: we obtain an expression just in terms of long-wavelength variables

$$\frac{\partial^2}{H^2} \Phi(\vec{x}, t) = \delta(\vec{x}, t) + \partial_i \partial_j Q_{ij}\, (\delta(\vec{x}, t), \ldots) + \ldots$$

This EFT is non-local in time

- For local EFT, we need hierarchy of scales.
  - In space we are ok





  - In time we are not ok: all modes evolve with time-scale of order Hubble

with Carrasco, Foreman and Green **1310**

Carroll, Leichenauer, Pollak **1310**

- $\Longrightarrow$ The EFT is local in space, non-local in time
  - Technically it does not affect much because the linear propagator is local in space

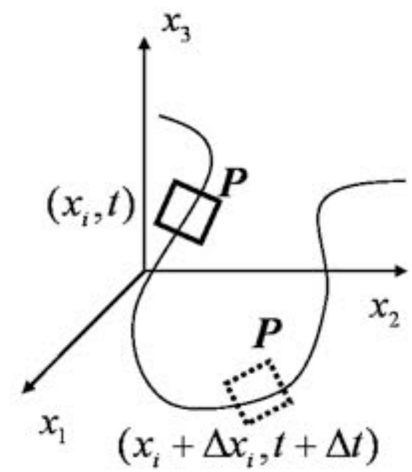
When do we stop?

- Similar to treatment for material polarizability: $\vec{d}_{\text{dipole}} \sim \alpha\, \vec{E}_{\text{electric}}$, $Q_{ij}^{\text{electric}} = c\, E_i E_j$, $\ldots$
- Short distance physics is taken into account by expectation value, response, and noise
- Poisson equation breaks when $\delta_{n,L}(\vec{x},\eta) \sim \partial_i \partial_j \mathcal{Q}^{ij}(\vec{x},\eta)$
  - gravitational potential from quadrupole moment ~ the one from center of mass
- By dimensional analysis, this happens for distances shorter than a critical length
  - the non-linear scale $k \gtrsim k_{\text{NL}}$
  - on long distances, $k \ll k_{\text{NL}}$, write as many terms as precision requires.
    - Manifestly convergent expansion in $\left(\dfrac{k}{k_{\text{NL}}}\right) \ll 1$

Connecting with the Eulerian Treatment

- In the universe, finite-size particles move

$$\vec{z}(\vec{q},t) = \vec{q} + \vec{s}(\vec{q},t)$$



- In Lagrangian space, we do not expand in $\vec{s}(\vec{q},t)$
- In Eulerian, we do: we describe particles from a fixed position
  - Expand in $k\, s \ll 1$
- There are three expansion parameters for a given wavenumber

$$\epsilon_{s>} = k^2 \int_k^\infty \frac{d^3k'}{(2\pi)^3} \frac{P_{11}(k')}{k'^2} \,,$$ Effect of Short Displacements

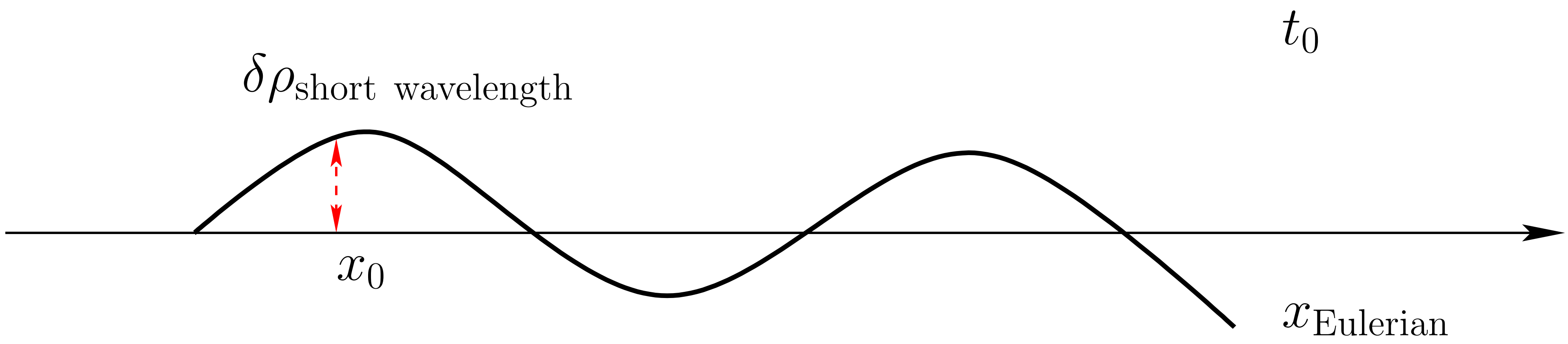
$$\epsilon_{\delta<} = \int_0^k \frac{d^3k'}{(2\pi)^3} P_{11}(k') \,,$$ Effect of Long Overdensities

$$\epsilon_{s<} = k^2 \int_0^k \frac{d^3k'}{(2\pi)^3} \frac{P_{11}(k')}{k'^2} \,,$$ Effect of Long Displacements: Lagrangian does not expands in this

The Effect of Long Displacements

$$\epsilon_{s<} = k^2 \int_0^k \frac{d^3k'}{(2\pi)^3} \frac{P_{11}(k')}{k'^2}$$

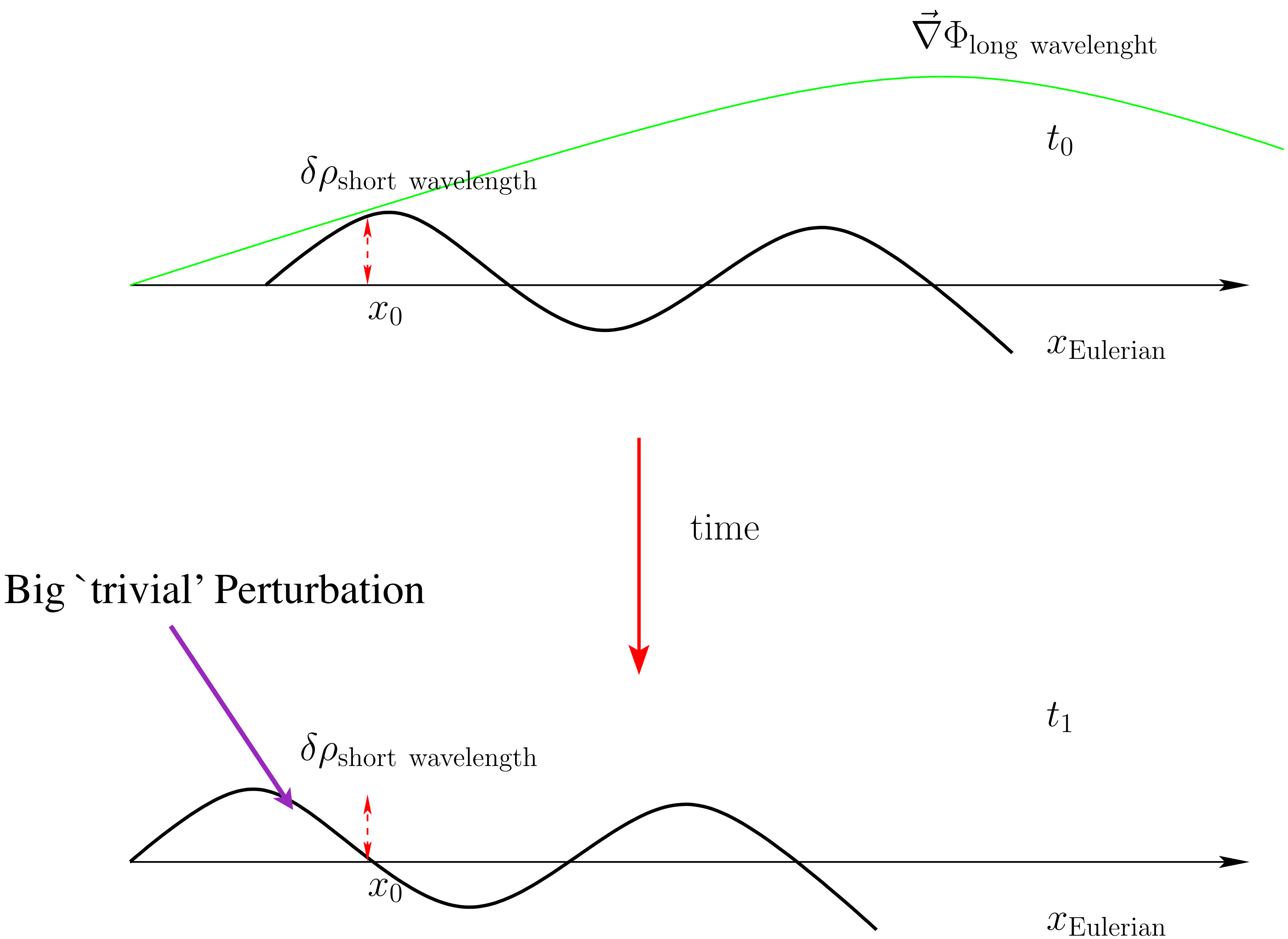
- Imagine a mode



The Effect of Long Displacements

- Add a long `trivial’ force (trivial by GR)
- Just Translation

$$\epsilon_{s<} = k^2 \int_0^k \frac{d^3k'}{(2\pi)^3} \frac{P_{11}(k')}{k'^2}$$

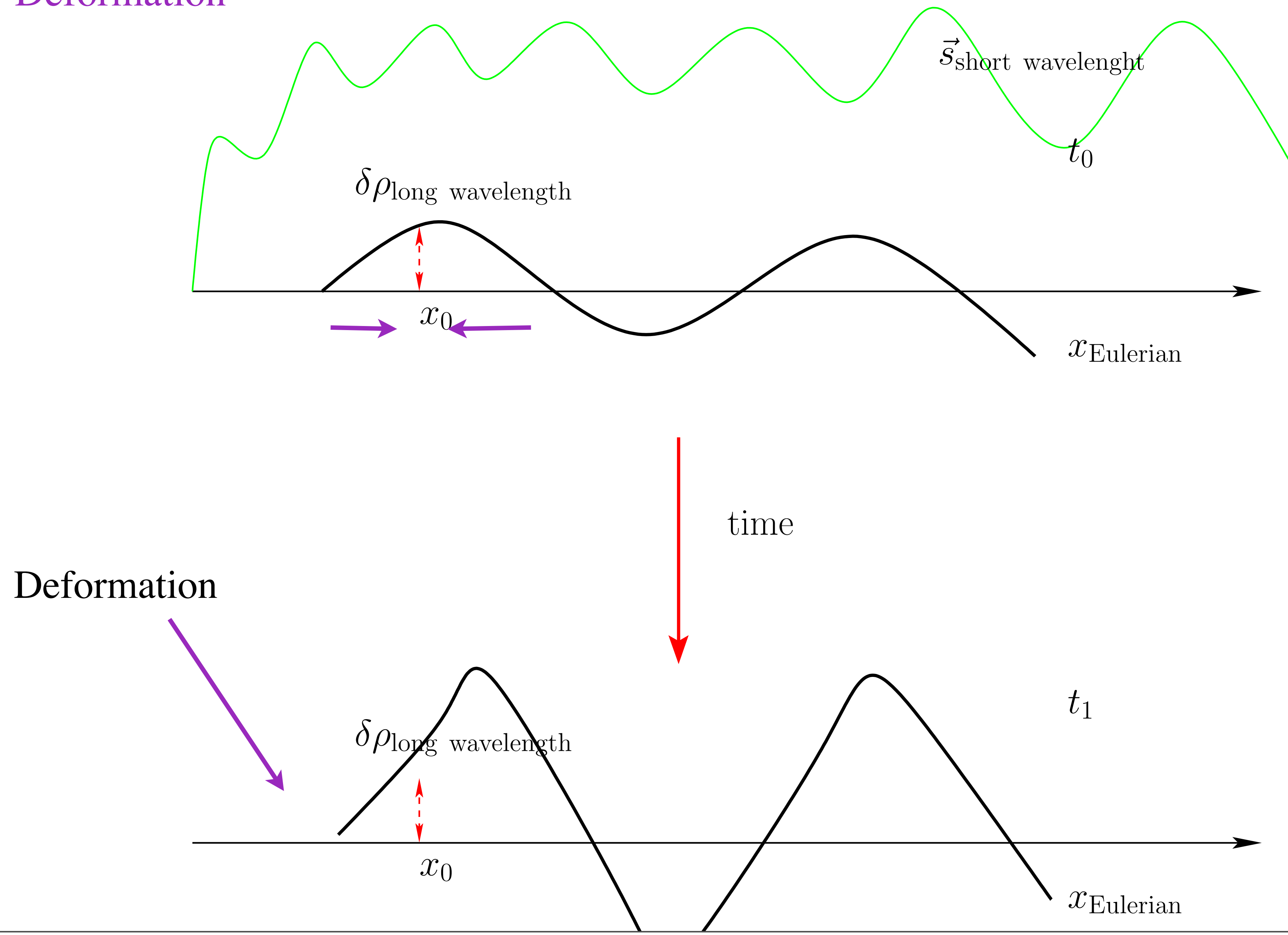


The Effect of Short Displacement

- Add a long `trivial' force (trivial by GR)

$$\epsilon_{s>} = k^2 \int_k^\infty \frac{d^3k'}{(2\pi)^3} \frac{P_{11}(k')}{k'^2}$$

- Deformation

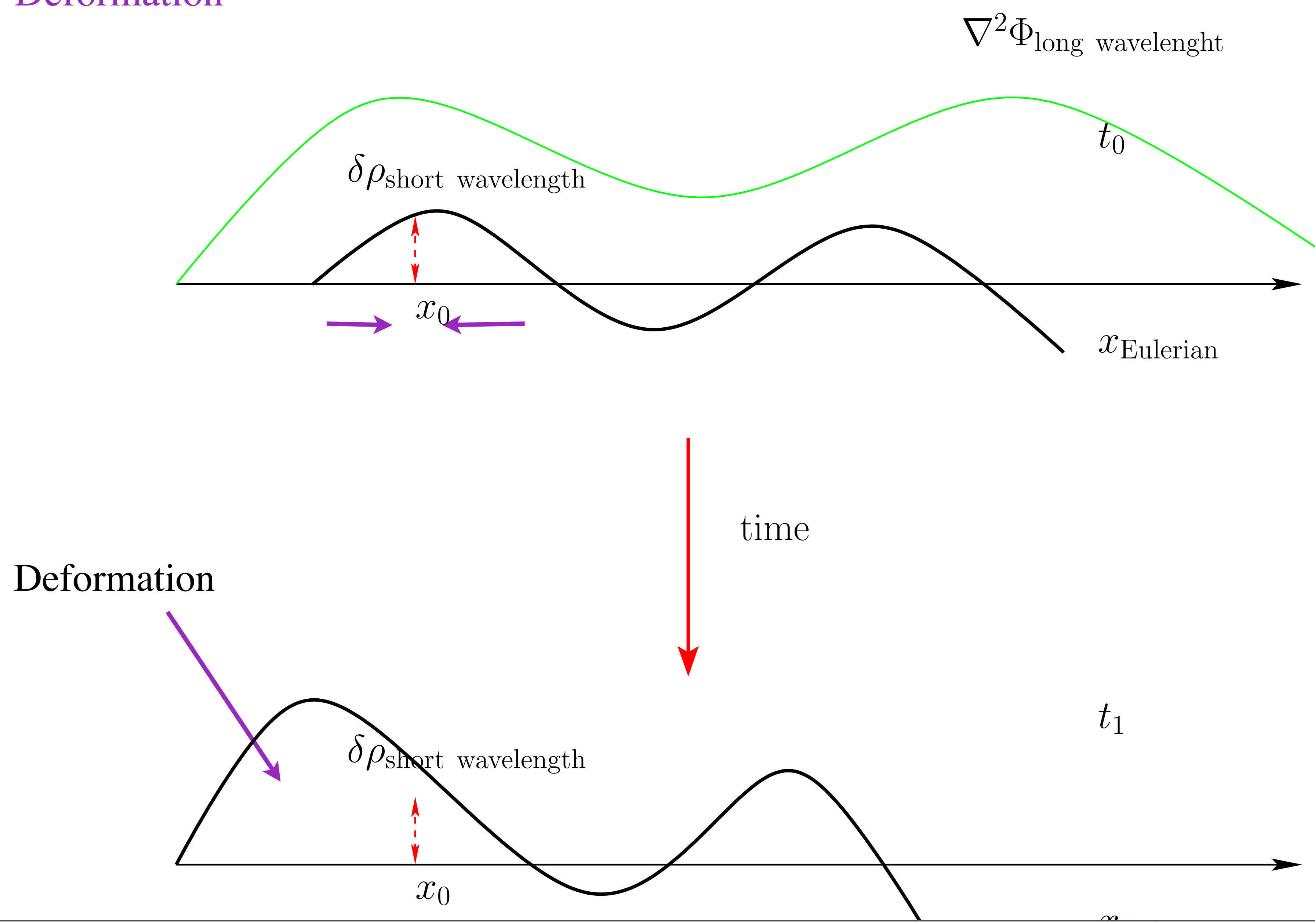


The Effect of Tidal Forces

- Add a long `trivial' force (trivial by GR)

$$\epsilon_{\delta<} = \int_0^k \frac{d^3k'}{(2\pi)^3} P_{11}(k') \ ,$$

- Deformation



Connecting with the Eulerian Treatment

- Expand in all parameters (Eulerian treatment)
- The resulting equations are equivalent to Eulerian fluid-like equations

$$\nabla^2\phi = H^2\frac{\delta\rho}{\rho}$$

$$\partial_t\rho + H\rho + \partial_i(\rho v^i) = 0$$

$$\dot{v}^i + Hv^i + v^j\partial_j v^i = \frac{1}{\rho}\partial_j\tau^{ij}$$

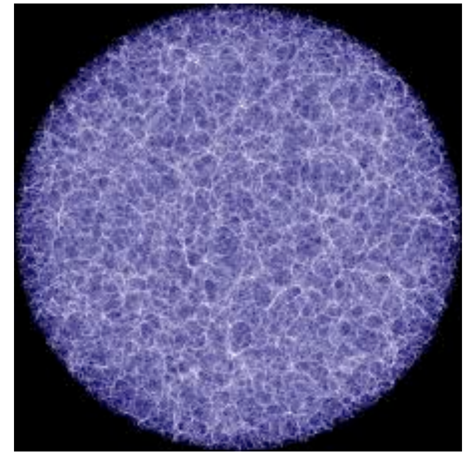
– here it appears a non trivial stress tensor for the long-distance fluid

$$\tau_{ij} = p_0\,\delta_{ij} + c_s^2\,\delta_{ij}\,\partial^2\delta\rho + \ldots$$

Perturbation Theory with the EFT

A non-renormalization theorem

- Can the short distance non-linearities change completely the overall expansion rate of the universe, possibly leading to acceleration without Λ ?



- In terms of the short distance perturbation, the effective stress tensor reads

$$\rho_L = \rho_S \left(1 + v_S^2 + \Phi_S\right)$$
$$p_L = \rho_S \left(2v_S^2 + \Phi_L\right)$$

- when objects virialize, the induced pressure vanish
  – ultraviolet modes do not contribute (like in SUSY)
- The backreaction is dominated by modes at the virialization scale

$\Rightarrow \quad w_{\text{induced}} \sim 10^{-5}$ with Baumann, Nicolis and Zaldarriaga **JCAP 2012**

–

Perturbation Theory within the EFT

- In the EFT we can solve iteratively (loop expansion) $\delta_\ell, v_\ell, \Phi_\ell \ll 1$

$$\nabla^2\phi = H^2\frac{\delta\rho}{\rho}$$

$$\partial_t\rho + H\rho + \partial_i(\rho v^i) = 0$$

$$\dot{v}^i + Hv^i + v^j\partial_j v^i = \frac{1}{\rho}\partial_j\tau^{ij}$$

$$\tau_{ij} = p_0\,\delta_{ij} + c_s^2\,\delta_{ij}\,\partial^2\delta\rho$$

Perturbation Theory within the EFT

- Regularization and renormalization of loops (scaling universe)
  - evaluate with cutoff. By dim analysis:

$$
\begin{aligned}
P_{1-\text{loop}} = {} & c_0^\Lambda \left(\frac{\Lambda}{k_{\rm NL}}\right)^2 \left(\frac{k}{k_{\rm NL}}\right) P_{11} + c_1^\Lambda \left(\frac{\Lambda}{k_{\rm NL}}\right) \left(\frac{k}{k_{\rm NL}}\right)^2 P_{11} \\
& + c_2^\Lambda \log\left(\frac{\Lambda}{k_{\rm NL}}\right) \left(\frac{k}{k_{\rm NL}}\right)^3 P_{11} + c_1^{\rm finite} \left(\frac{k}{k_{\rm NL}}\right)^3 P_{11} + \text{subleading in } \frac{k}{k_{\rm NL}}
\end{aligned}
$$

Perturbation Theory within the EFT

- Regularization and renormalization of loops (scaling universe)
  - evaluate with cutoff. By dim analysis:

$$P_{1-\text{loop}} = \cancel{c_0^\Lambda \left(\frac{\Lambda}{k_{\rm NL}}\right)^2 \left(\frac{k}{k_{\rm NL}}\right)} P_{11} + c_1^\Lambda \left(\frac{\Lambda}{k_{\rm NL}}\right) \left(\frac{k}{k_{\rm NL}}\right)^2 P_{11}$$
$$+ c_2^\Lambda \log \cancel{\left(\frac{\Lambda}{k_{\rm NL}}\right) \left(\frac{k}{k_{\rm NL}}\right)}^3 P_{11} + c_1^{\rm finite} \left(\frac{k}{k_{\rm NL}}\right)^3 P_{11} + \text{subleading in } \frac{k}{k_{\rm NL}}$$

  - absence of counterterm $\qquad \tau_{ij} = p_0\, \delta_{ij} + c_s^2\, \delta_{ij}\, \partial^2 \delta\rho$

Perturbation Theory within the EFT

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$$P_{1-\text{loop}} = \cancel{c_0^\Lambda \left(\frac{\Lambda}{k_{\rm NL}}\right)^2 \left(\frac{k}{k_{\rm NL}}\right)} P_{11} + c_1^\Lambda \left(\frac{\Lambda}{k_{\rm NL}}\right) \left(\frac{k}{k_{\rm NL}}\right)^2 P_{11}$$
$$+ \cancel{c_2^\Lambda \log\left(\frac{\Lambda}{k_{\rm NL}}\right)\left(\frac{k}{k_{\rm NL}}\right)^3} P_{11} + c_1^{\rm finite}\left(\frac{k}{k_{\rm NL}}\right)^3 P_{11} + \text{subleading in } \frac{k}{k_{\rm NL}}\,\frac{k}{k_{\rm NL}}$$

  - absence of counterterm $\qquad \tau_{ij} = p_0\,\delta_{ij} + c_s^2\,\delta_{ij}\,\partial^2\delta\rho$

$$\Longrightarrow \qquad P_{1-\text{loop},\ \text{counter}} = c^\Lambda_{\rm counter}\left(\frac{k}{k_{\rm NL}}\right)^2 P_{11}$$

$$\Longrightarrow \qquad c^\Lambda_{\rm counter} = -c_1^\Lambda + \delta c_{\rm counter}\left(\frac{k_{\rm NL}}{\Lambda}\right)$$

$$\Longrightarrow$$
$$P_{1-\text{loop}} + P_{1-\text{loop},\ \text{counter}} = \delta c_{\rm counter}\left(\frac{k}{k_{\rm NL}}\right)^2 P_{11} + c_1^{\rm finite}\left(\frac{k}{k_{\rm NL}}\right)^3 P_{11}$$

Calculable terms in the EFT

- Has everything being lost?

$$P_{1-\text{loop}} + P_{1-\text{loop},\ \text{counter}} = \delta c_{\text{counter}} \left(\frac{k}{k_{\text{NL}}}\right)^2 P_{11} + c_1^{\text{finite}} \left(\frac{k}{k_{\text{NL}}}\right)^3 P_{11}$$

  – to make result finite, we need to add a counterterm with finite part

    - need to fit to data (like a coupling constant), but cannot fit the k-shape

Calculable terms in the EFT

- Has everything being lost?

$$P_{1-\text{loop}} + P_{1-\text{loop, counter}} = \delta c_{\text{counter}} \left(\frac{k}{k_{\text{NL}}}\right)^2 P_{11} + c_1^{\text{finite}} \left(\frac{k}{k_{\text{NL}}}\right)^3 P_{11}$$

  – to make result finite, we need to add a counterterm with finite part
    - need to fit to data (like a coupling constant), but cannot fit the k-shape

  – the subleading finite term is not degenerate with a counterterm.
    - it cannot be changed
    - it is calculable by the EFT
      – so it predicts an observation $c_1^{\text{finite}} = 0.044$

Lesson from Renormalization

- Each loop-order L contributed a finite, calculable term of order

$$P_{\mathrm{L-loops}} \sim \left(\frac{k}{k_{\mathrm{NL}}}\right)^{L}$$

  – each higher-loop is smaller and smaller

- This happens after canceling the divergencies with counterterms

$$P_{\mathrm{L-loops;\ without\ counterterms}} = \left(\frac{\Lambda}{k_{\mathrm{NL}}}\right)^{L} \frac{k^2}{k_{\mathrm{NL}}^2} P(k)$$

  - each loop contributes the same

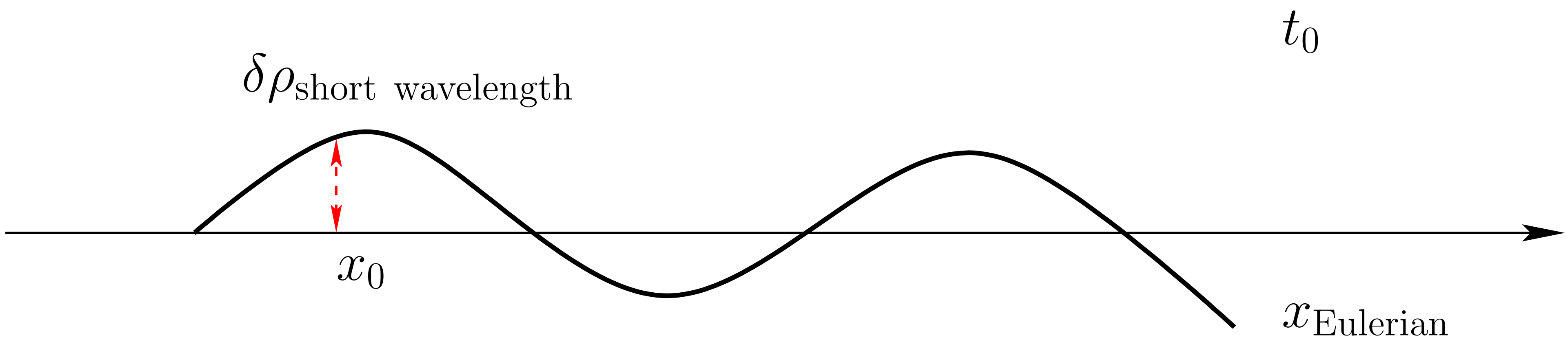
- Up to 2-loops, we need only the 1-loop counterterm

IR-resummation

with Zaldarriaga **1404**

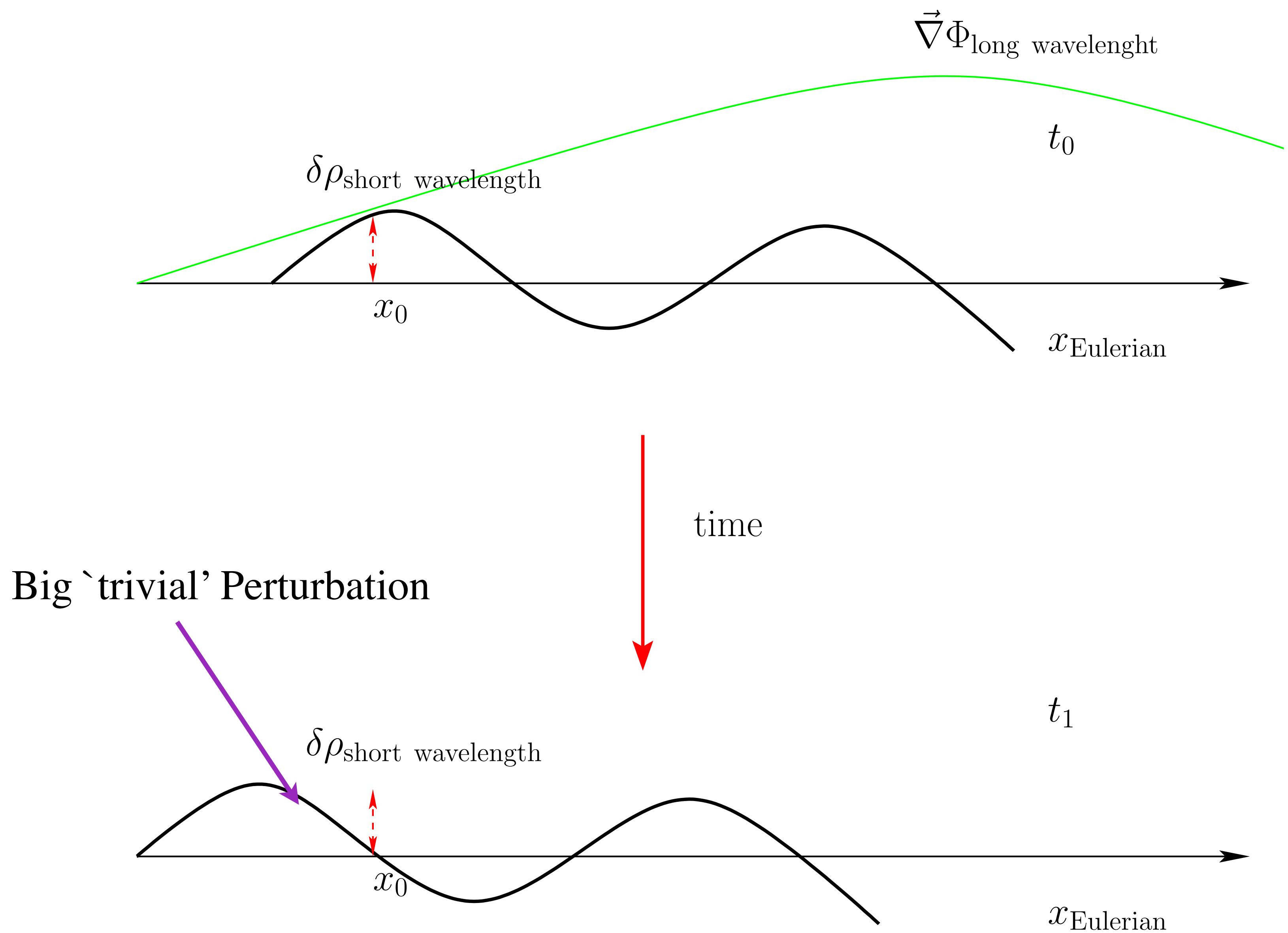
The Effect of Long-modes on Shorter ones

- In Eulerian treatment



The Effect of Long-modes

- Add a long `trivial' force (trivial by GR)
- This tells you that one can resume the IR modes: this is the Lagrangian treatment



The Effect of Long-modes

- Two effects

$$\vec{\pi}(\vec{x}) \quad \longrightarrow \quad \vec{\pi}_{\rm inertial}(\vec{\tilde{x}}) = \vec{\pi}(\vec{x}(\vec{\tilde{x}})) + \rho(\vec{\tilde{x}})\ \vec{v}(\vec{\tilde{x}})$$

  - Shift in coordinates
  - Shift in field

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The Effect of Long-modes

- Two effects

$$\vec{\pi}(\vec{x}) \quad \to \quad \vec{\pi}_{\text{inertial}}(\tilde{\vec{x}}) = \vec{\pi}(\vec{x}(\tilde{\vec{x}})) + \rho(\tilde{\vec{x}})\, \vec{v}(\tilde{\vec{x}})$$

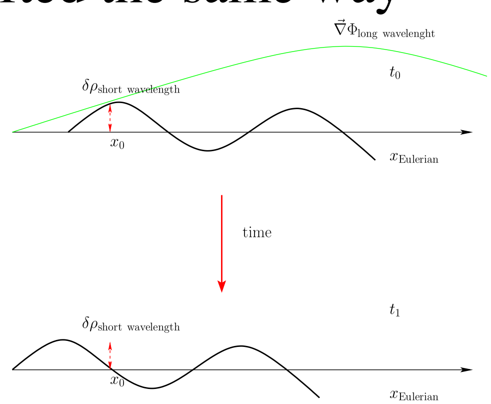
  – Shift in coordinates

  – Shift in field

- For fields that are scalar, this naively implies, by GR, that there are no IR effects in Fourier space at equal time correlators

  – both modes are shifted the same way

with Frieman and Scoccimarro **1996**

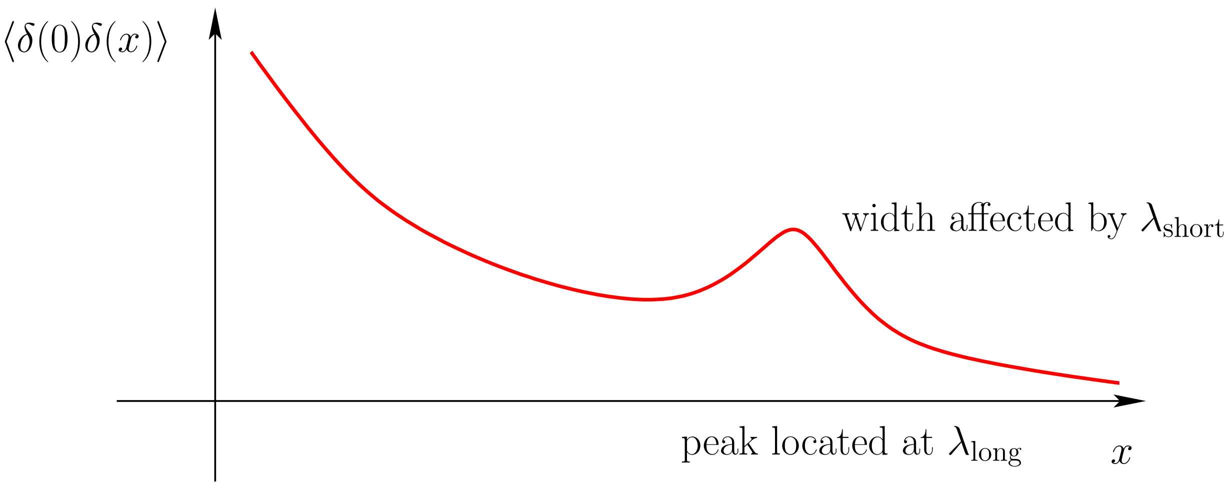


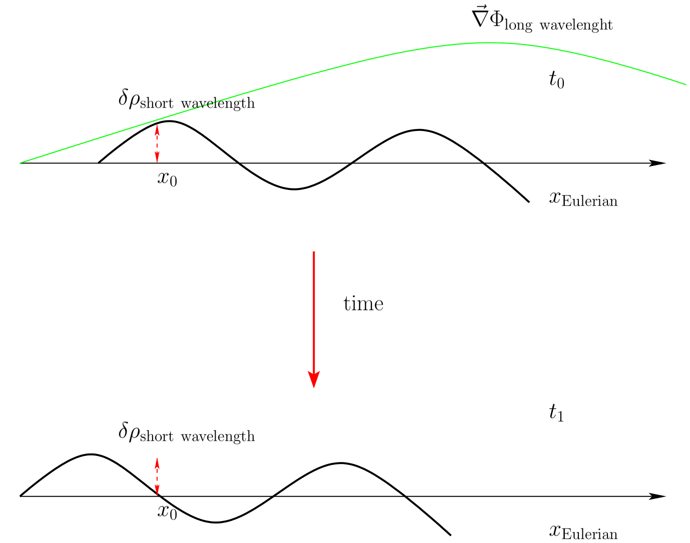
with Carrasco, Foreman and Green **1304**
used to find the so-called consistency conditions in GR

Creminelli, Norena, Simonovic **1309**

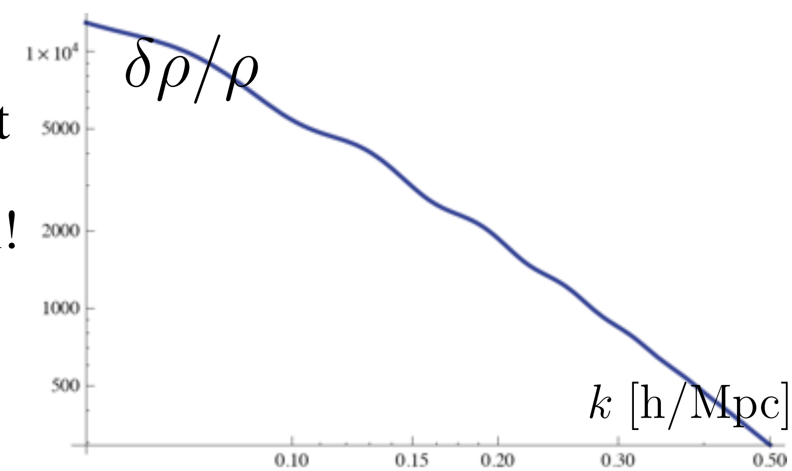
The Effect of Long-modes

with Zaldarriaga **1304**



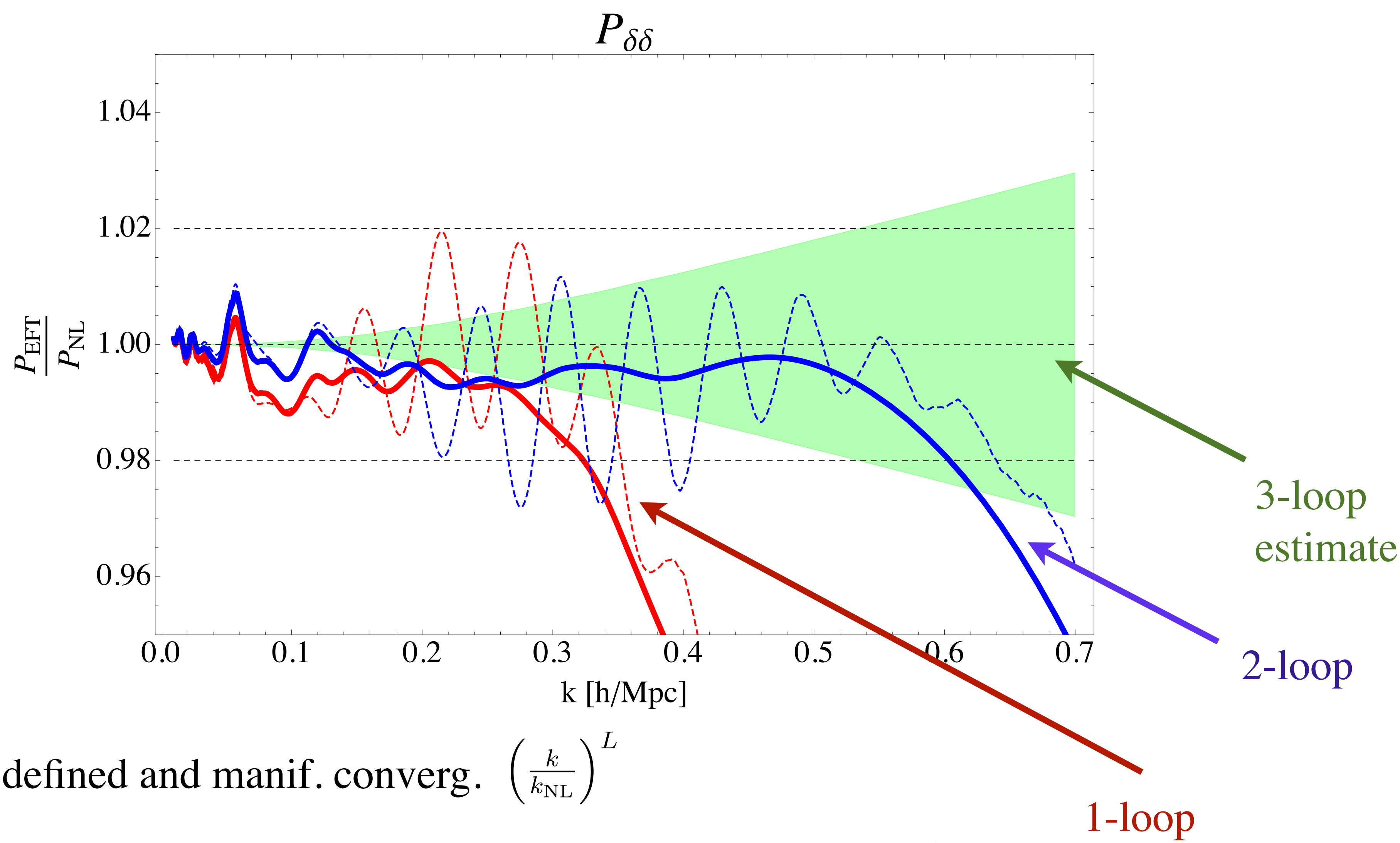


- The universe has features!
- Even on equal time correlators, IR modes of order the BAO scale do not cancel!
  - In Fourier space these are the wiggles
- To compute the width, IR-BAO modes are relevant
- But they just do kinematics, so we can resum them!



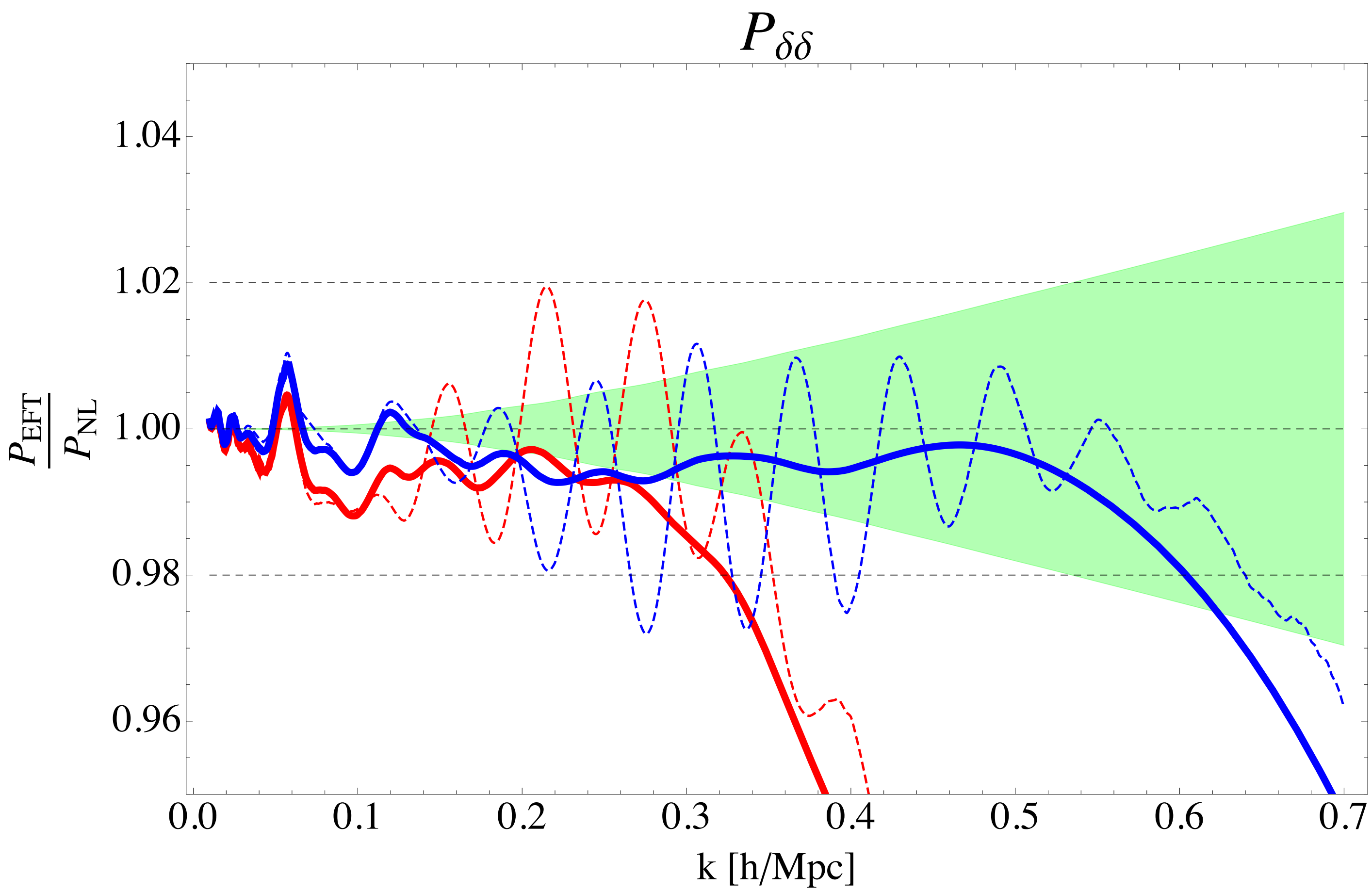
Results

EFT of Large Scale Structures



- Well defined and manif. converg. $\left(\frac{k}{k_{\mathrm{NL}}}\right)^L$
- Every perturbative order improves the agreement as it should
- We know when we should fail, and we fail when we should

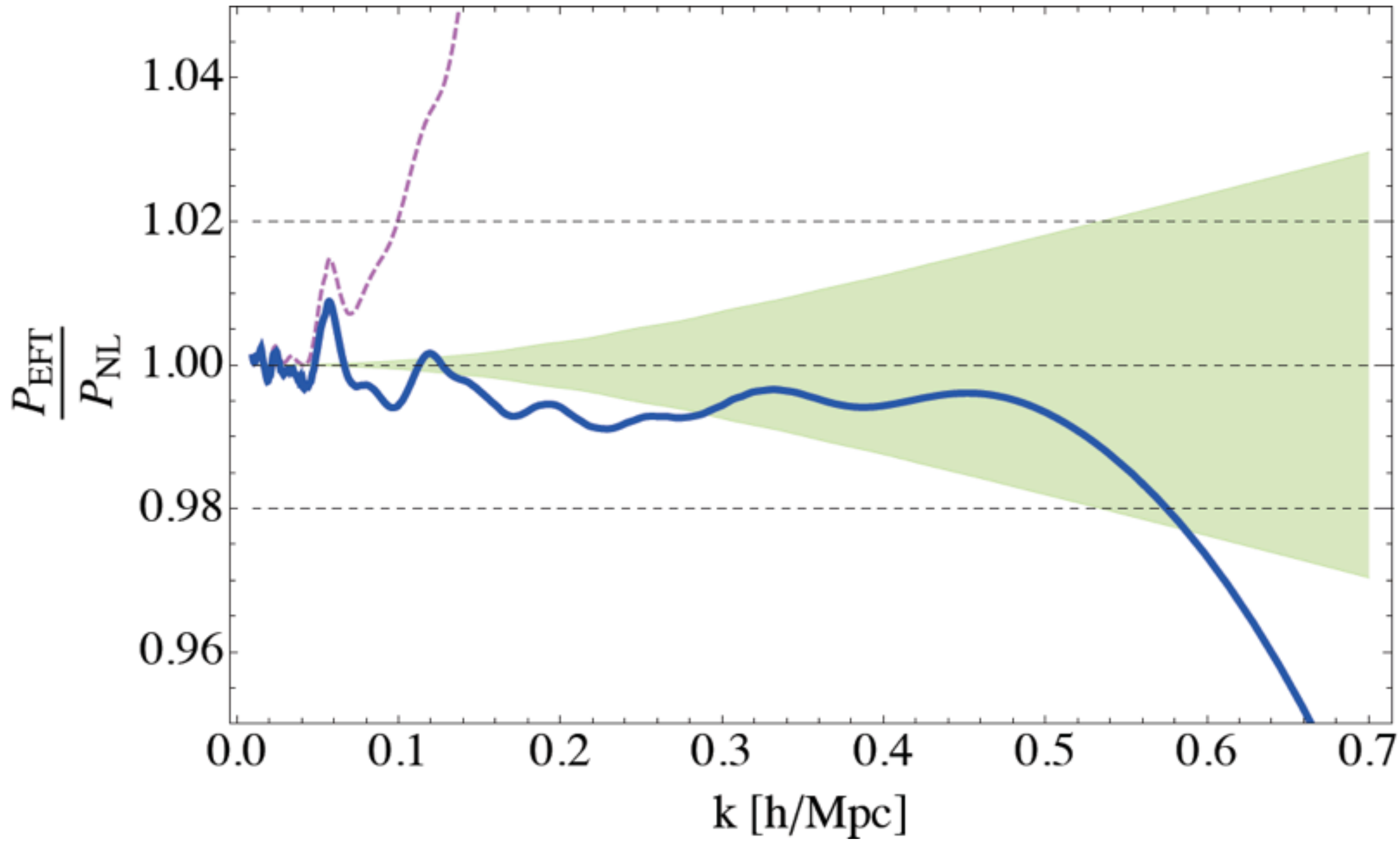
EFT of Large Scale Structures



- The lines with oscillations are obtained without resummation in the IR
  - Getting the BAO peak wrong

with Carrasco, Foreman and Green **1310**

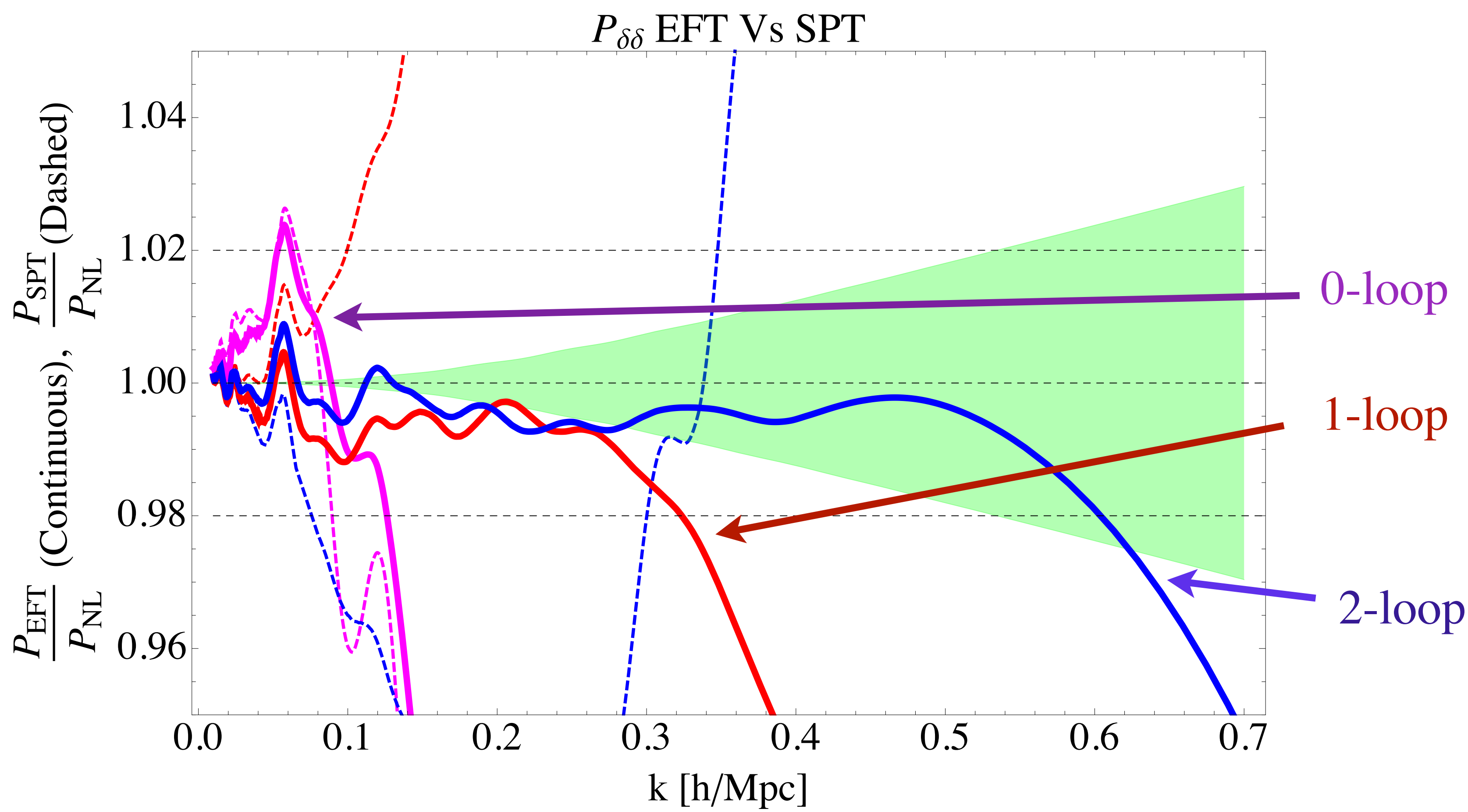
EFT of Large Scale Structures



- we fit until $k_{\rm max} \simeq 0.6\, h\, {\rm Mpc}^{-1}$, as where we should stop fitting
  - – there are 200 more quasi linear modes than previously believed!

with Zaldarriaga **1404**

EFT of Large Scale Structures



- Comparison with Standard Treatment (feel free to ask about RPT)
- For the EFT, change from 1-loop to 2-loop predicted

$$P_{\text{EFT-2-loop}} = P_{11} + P_{\text{1-loop}} + P_{\text{2-loop}} - 2\,(2\pi)(c^2_{s(1)} + c^2_{s(2)})\frac{k^2}{k^2_{\rm NL}}P_{11} + (2\pi)c^2_{s(1)}P^{(c_s,p)}_{\text{1-loop}} + (2\pi)^2 c^4_{s(1)}\frac{k^4}{k^4_{\rm NL}}P_{11}$$

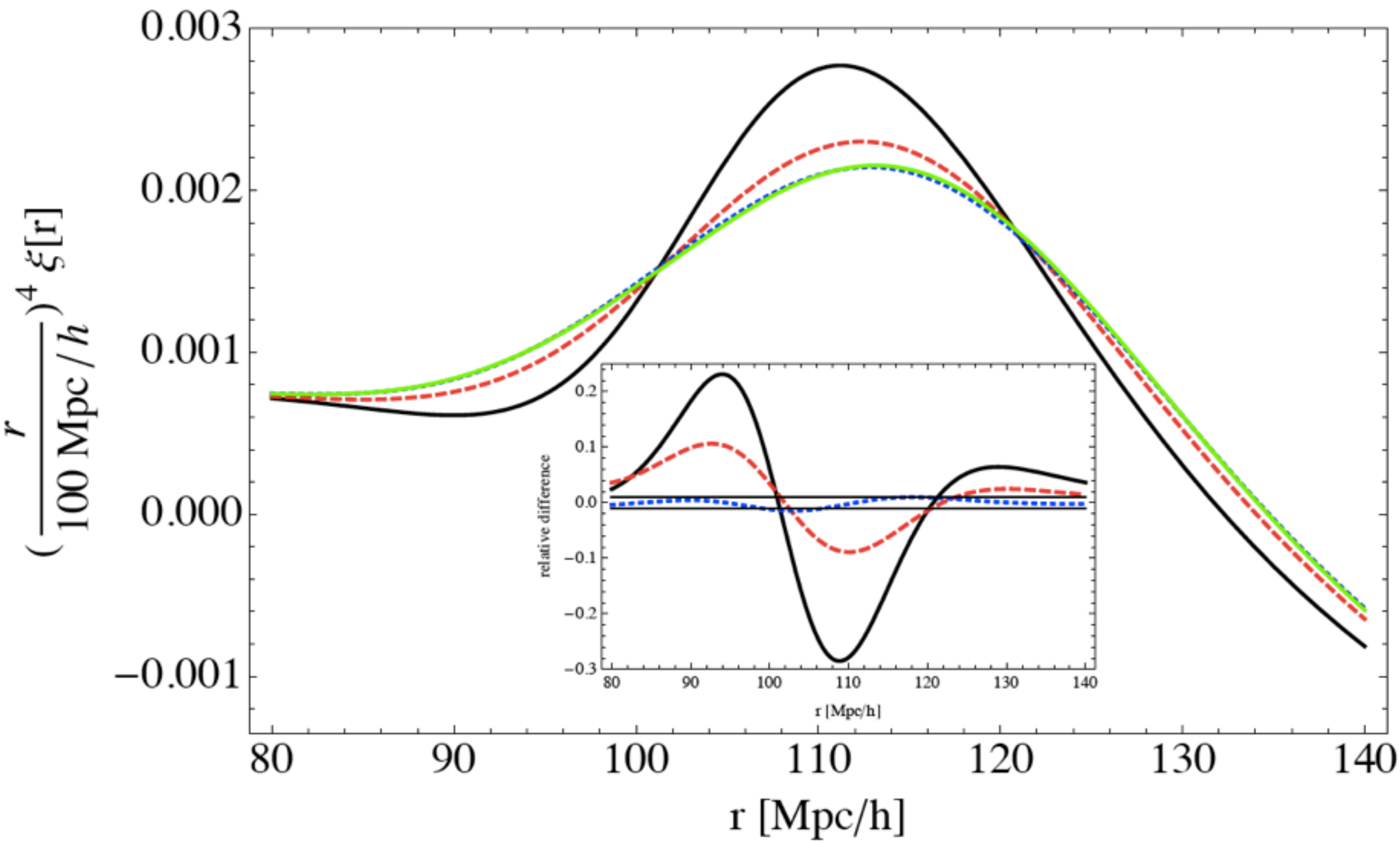
– the other new terms are clearly important

– they `conspire' to the right answer

The BAO peak in `5 minutes'

- The IR-resummation is crucial to get the BAO peak right.
  - we can do this very quickly.

with Zaldarriaga **1404**

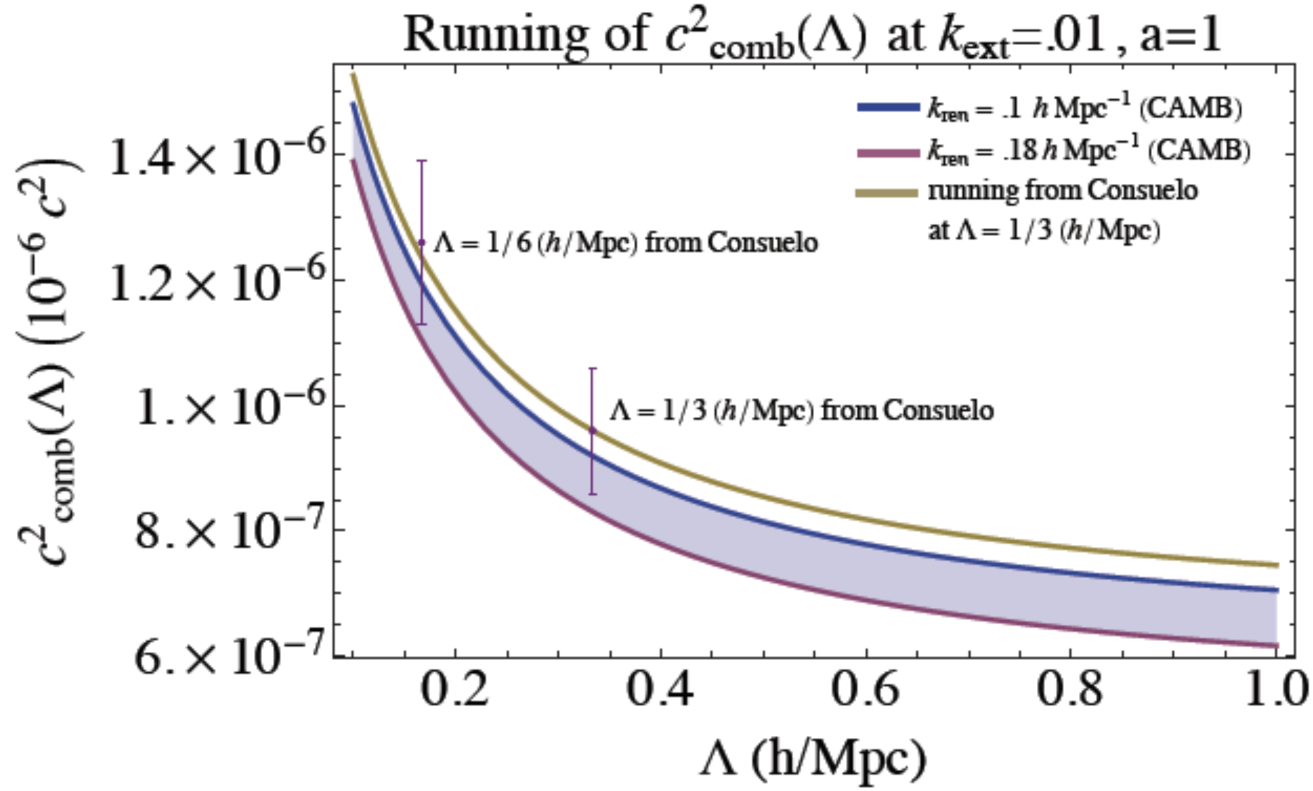


Measuring Parameters from small N-body Simulations

Measuring parameters from N-body sims.

- The EFT parameters can be measured from small N-body simulations
  - similar to what happens in QCD: lattice sims
- As you change smoothing scale, the result changes

$$\frac{d\,c_s}{d\Lambda} = \frac{d}{d\,\Lambda}\int^{\Lambda} d^3k\ P_{13}(k)$$



- Perfect agreement with fitting at low energies
  - like measuring F_π from lattice sims and $\pi\pi$ scattering

with Carrasco and Hertzberg **JHEP 2012**

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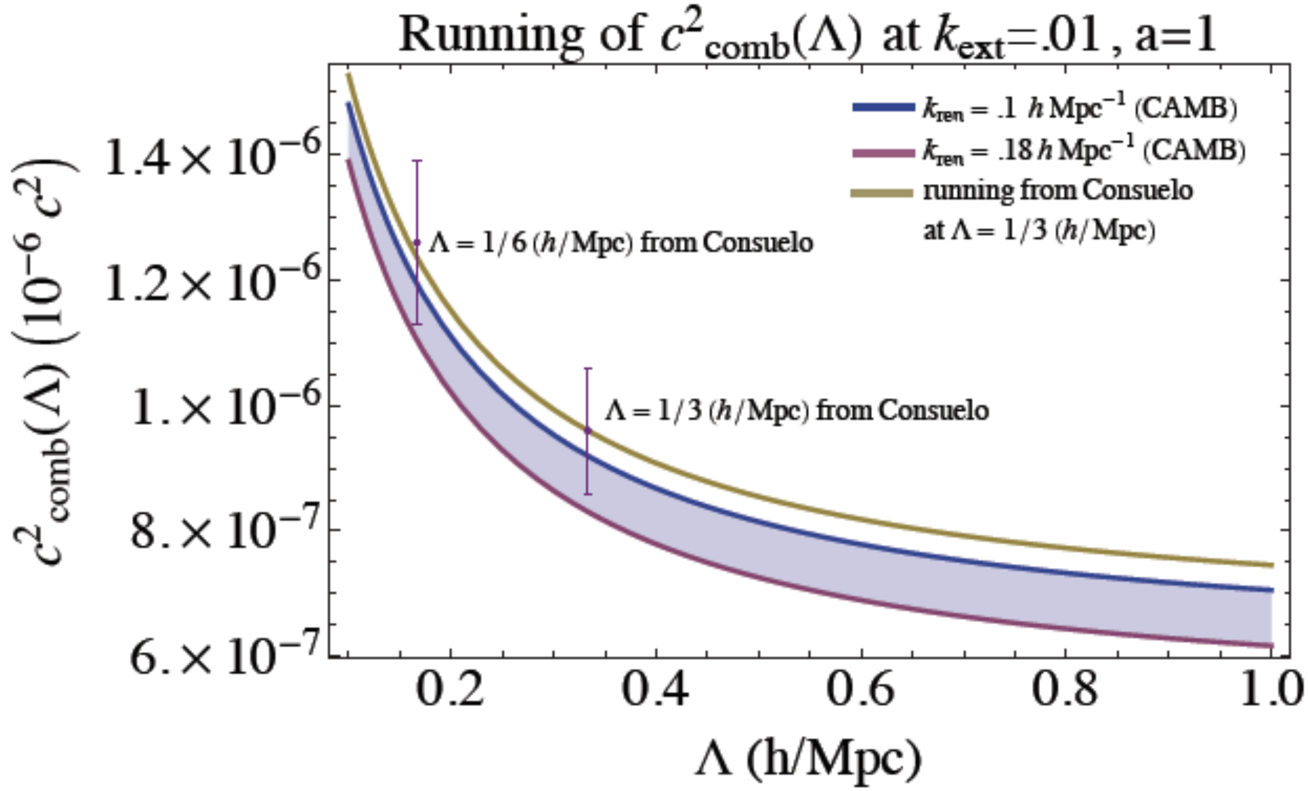

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$$[\partial_i\partial_j v_k](\vec r) = [\partial_i\partial_j \pi_k](\vec r)/[\rho](\vec r) - [\partial_i \pi_k](\vec r)[\partial_j\rho]/([\rho](\vec r))^2 - [\partial_j \pi_k](\vec r)[\partial_i\rho]/([\rho](\vec r))^2$$
$$- [\pi_k](\vec r)[\partial_i\partial_j\rho](\vec r)/([\rho](\vec r))^2 + 2[\pi_k](\vec r)[\partial_i\rho](\vec r)[\partial_j\rho](\vec r)/([\rho](\vec r))^3$$

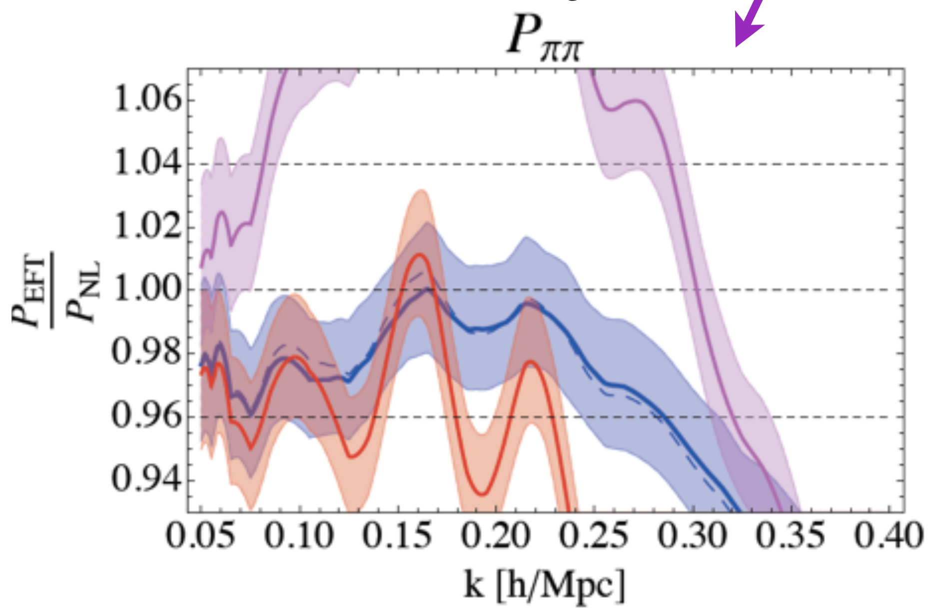
  - UV dof

ırrasco and Hertzberg **JHEP 2012**

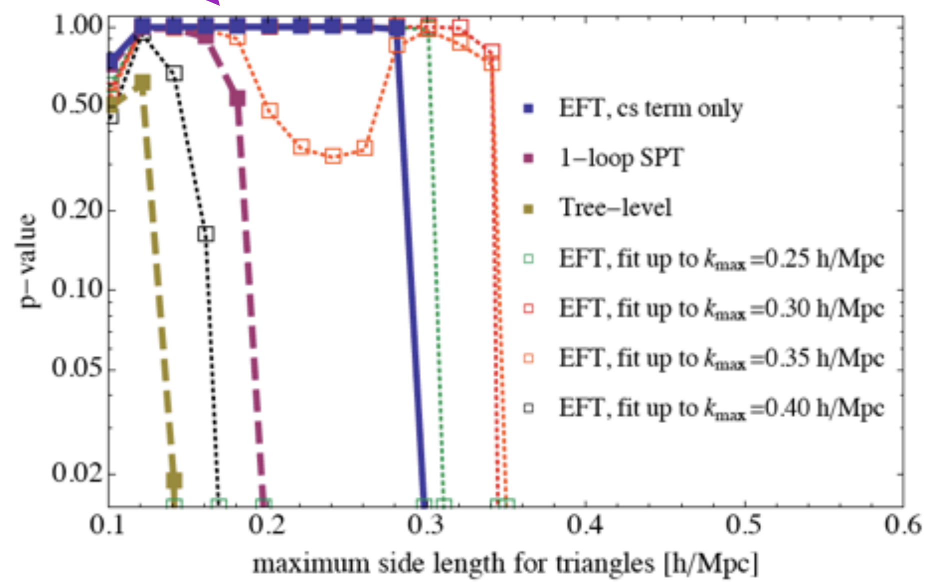
Other Observables

Momentum and Bispectrum

with Zaldarriaga **1404**

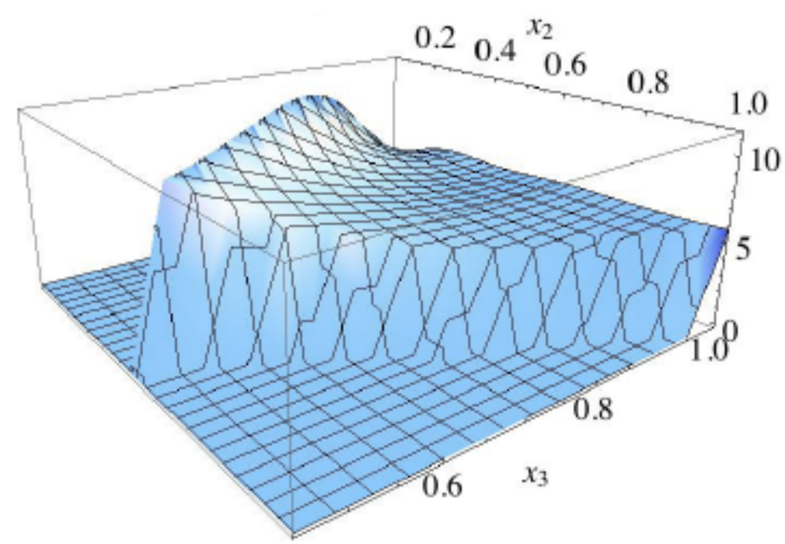


with Angulo, Foreman and Schmittful **1406**



- At one-loop, similarly great results
  - with no additional parameter
  - as good as they should
  - very non-trivial functional forms
- Similar formulas just worked out for Bias

Senatore **1406** See also (McDoland and Roy **0902**)



  - and Redshfit space distortions

with Zaldarriaga **1409**

Velocity field

- Momentum is a natural quantity, as connected to density by conservation law
- Velocity is not a natural quantity $\vec{v}(\vec{x}) = \frac{\vec{\pi}(\vec{x})}{\rho(\vec{x})}$
- It is a local composite operator: needs its own new counterterms:

$$v_{l,R}(\vec{x},t) = v_l(\vec{x},t) - e_1 \partial\delta(\vec{x},t) + \cdots$$

with Carrasco, Foreman and Green **1310**

  – no new counterterm for the equations

- Because of this, and because it is a viscous fluid, we generate vorticity

$$\langle \omega_k^2 \rangle \sim \alpha_1 \left(\frac{k}{k_{\rm implement.}} \right)^2 + \alpha_2 \left(\frac{k}{k_{\rm NL}} \right)^{\sim 3}$$

  – from local counterterm

  – from viscosity

- Predicted result seems to be verified in sims

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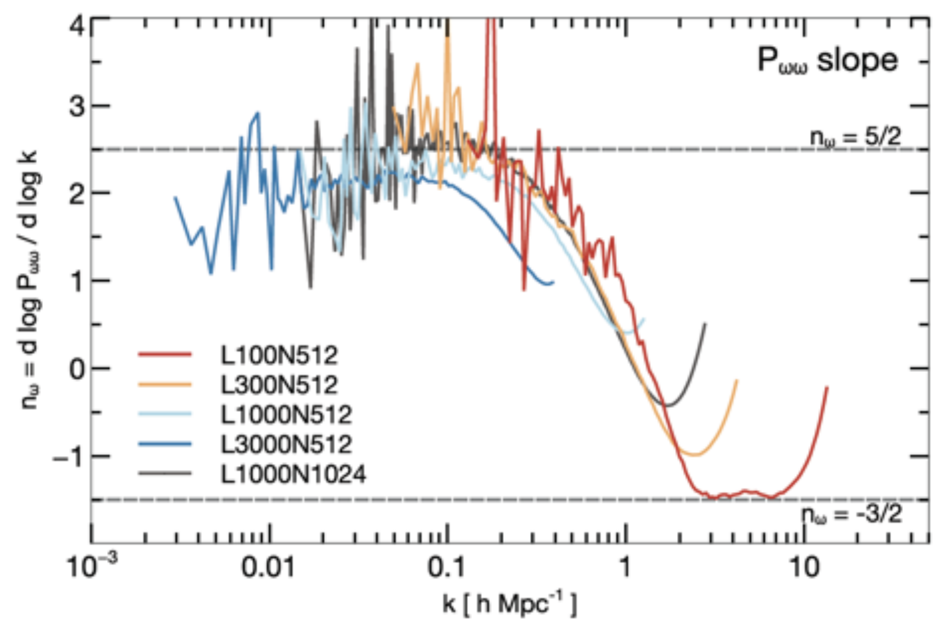
Hahn, Angulo, Abel, **to appear**

see also Pueblas and Scoccimarro **08**



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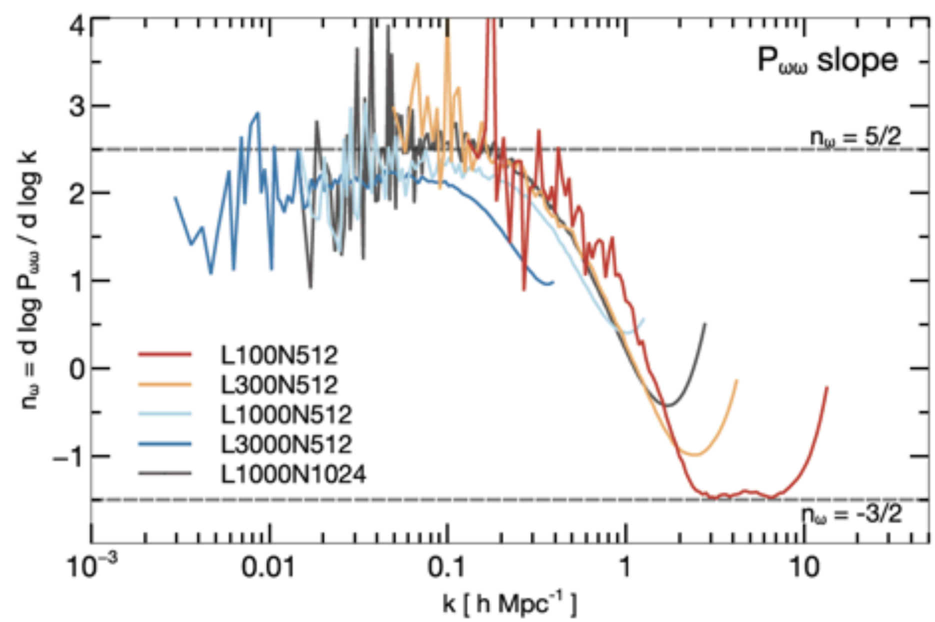
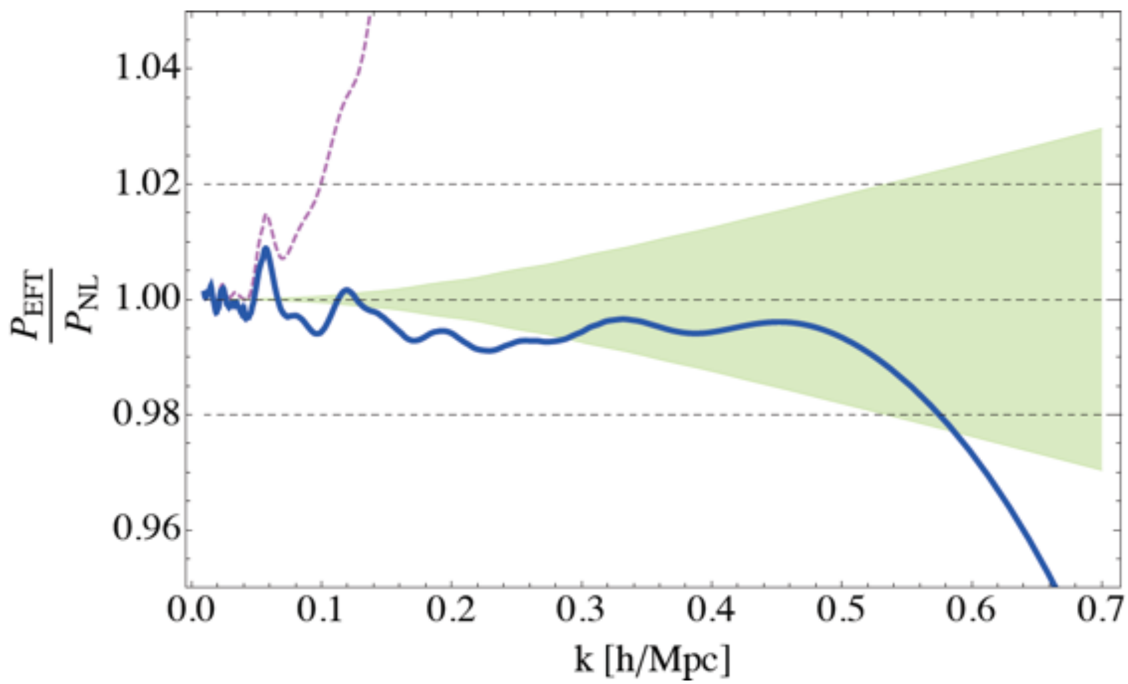


- Predicted result seems to be verified in sims
- Former analytic techniques got zero

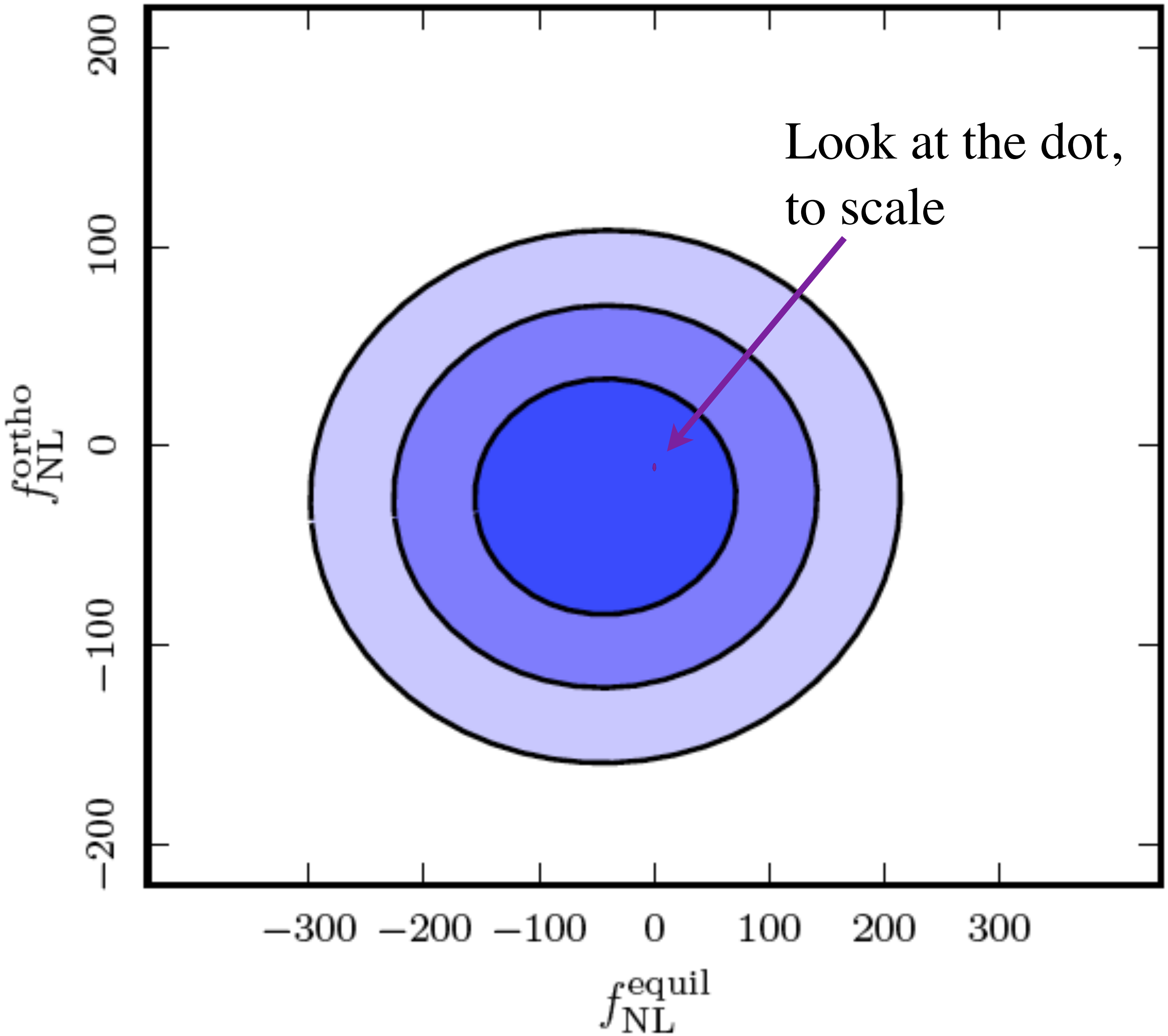
End to SPT-like resummations

EFT of Large Scale Structures



- A manifestly convergent perturbation theory $\left(\frac{k}{k_{\rm NL}}\right)^L$
- we fit until $k_{\rm max} \simeq 0.6\, h\, {\rm Mpc}^{-1}$, as where we should stop fitting
  - there are 200 more quasi linear modes than previously believed!
  - huge impact on possibilities, for ex: $f_{\rm NL}^{\rm equil.,\, orthog.} \lesssim 1$
- Can all of us handle it?! This is an huge opportunity and a challenge for us

With this


Look at the dot,
to scale
f_{NL}^{ortho}
200
100
0
−100
−200
−300
−200
−100
0
100
200
300
f_{NL}^{equil}

Conclusions

- Many (most?) of the features of QFT appear in the EFT of LSS:
  - Loops, divergencies, counterterms and renormalization
  - non-renormalization theorems
  - Calculable and non-calculable terms
  - Measurements in lattice and lattice-running
  - IR-divergencies
- Results seem to be amazing, many calculations and verifications to do:
  - like if we just learned perturbative QCD, and LHC was soon turning on
    - higher n-point functions
    - Validation with simulation
  - With a growing number of groups (Caltech, Princeton, IAS, Cambridge, CEA, Zurich..., just after 2-loop result, a workshop was organized by Princeton)
- If this works, the 10-yr future of Early Cosmology is good, even with no luck